









# Transport of High-energy Charged Particles through Spatially Intermittent Turbulent Magnetic Fields

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## Abstract

Identifying the sources of the highest energy cosmic rays requires understanding how they are deflected by the stochastic, spatially intermittent intergalactic magnetic field. Here we report measurements of energetic charged-particle propagation through a laser-produced magnetized plasma with these properties. We characterize the diffusive transport of the particles experimentally. The results show that the transport is diffusive and that, for the regime of interest for the highest energy cosmic rays, the diffusion coefficient is unaffected by the spatial intermittency of the magnetic field.

*Unified Astronomy Thesaurus concepts:* [High energy astrophysics \(739\)](#); [Ultra-high-energy cosmic radiation \(1733\)](#); [Magnetic fields \(994\)](#); [Laboratory astrophysics \(2004\)](#); [Plasma astrophysics \(1261\)](#); [Intergalactic medium \(813\)](#); [Particle astrophysics \(96\)](#)

## 1. Introduction

The interplay between charged particles and stochastic magnetic fields generated by plasma turbulence is crucial to understanding how cosmic rays propagate through space (Strong et al. 2007; Zweibel 2013; Schlickeiser 2015). A key parameter for determining the underlying nature of charged-particle diffusion is the ratio of the particle gyroradius  $r_g$  to the correlation length  $\ell_B$  of the magnetic turbulence. For the vast majority of cosmic rays detected at the Earth, this ratio is small. These are particles that are well confined by the Galactic magnetic field. But for cosmic rays more energetic than about 10 EeV, the ratio is larger than unity. These ultrahigh-energy cosmic rays (UHECRs) are not confined to the Milky Way and are presumed to be extragalactic in origin. Identifying their sources requires understanding how they are deflected by the intergalactic magnetic field, which appears to be stochastic and spatially intermittent. Recent data from the Parker Solar Probe mission have also indicated the presence of non-Gaussian magnetic fields near the Sun (Bandyopadhyay et al. 2020).

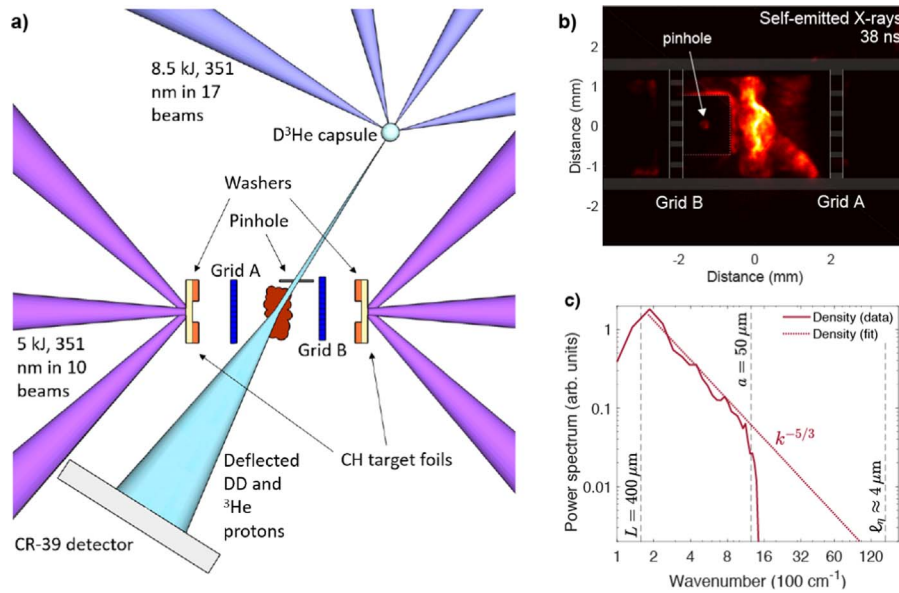
To study the propagation of cosmic rays, a theoretical framework has been developed based on direct numerical simulations of particle trajectories (e.g., Sigl et al. 2003, 2004) and statistical techniques (see Shalchi 2009 for a review). In particular, it has been shown (Jokipii 1966) that random small-amplitude fluctuations of the magnetic field superimposed on a mean background field lead to diffusive particle propagation. As a result, standard (Markovian) diffusion is widely used in

modeling cosmic-ray transport (e.g., Globus et al. 2008, 2019; Kotera & Lemoine 2008; Globus & Piran 2017), although anomalous diffusion has been shown to occur in special cases (Jokipii & Parker 1969; Reville et al. 2008; Lazarian & Yan 2014), including resonant scattering of charged particles in spatially intermittent magnetic fields (Shukurov et al. 2017).

Past laboratory experiments have studied particle transport in diffuse plasmas with strong mean magnetic fields (Gustafson et al. 2012; Anderson et al. 2013; Bovet et al. 2015; Furno et al. 2015), but the regime that is relevant to UHECR transport in the intergalactic medium (IGM), i.e., a stochastic, spatially intermittent magnetic field with zero mean ( $\langle \rangle = 0$ ), and under conditions of weak magnetization ( $r_g \gg \ell_B$ ), has not been studied theoretically, numerically, or experimentally.

## 2. Laser-driven Experiments

Here we report the results of laboratory experiments that focus on this regime. We carried out these experiments at the Omega Laser Facility at the Laboratory for Laser Energetics at the University of Rochester (Boehly et al. 1997). A high-velocity, magnetized, turbulent plasma was generated (Figure 1), employing the same platform as previously used to demonstrate dynamo amplification of magnetic fields (Tzeferacos et al. 2017, 2018). Three-dimensional simulations with the radiation-magnetohydrodynamics code FLASH (Fryxell et al. 2000; Dubey et al. 2009; Tzeferacos et al. 2015) guided and informed the experimental design, including



**Figure 1.** Experimental setup. (a) Schematic of the experimental platform, showing the target components and the configuration of the proton radiography experimental diagnostic. (b) X-ray self-emission from the interaction region at  $t = 38$  ns after the start of the laser drive. (c) Power spectrum of the density fluctuations recovered from the fluctuations of X-ray self-emission. The resolution of the diagnostic (the size of the pinhole employed on the X-ray framing camera) is  $50 \mu\text{m}$ , which is below the driving scale, but above the dissipative scales of the plasma.

target specifications and the timing of diagnostics (Tzeferacos et al. 2017).

The platform consists of two  $50 \mu\text{m}$  thick polystyrene (CH) foils attached to a pair of  $230 \mu\text{m}$  thick CH washers, with  $400 \mu\text{m}$  diameter machined “wells” that act as collimators, placed  $8 \text{ mm}$  apart. Between the two targets we position a pair of grids, comprised of periodic  $300 \mu\text{m}$  holes and  $100 \mu\text{m}$  wires, placed  $4 \text{ mm}$  apart. The grid patterns are shifted to break the mirror symmetry of the system. Each foil is irradiated with  $5 \text{ kJ}$  of energy during a  $10 \text{ ns}$  pulse (10 frequency-tripled laser beams on each foil, staggered in time). The drive produces two counterpropagating plasma flows, which pass through the pair of grids, meet, shear each other, and become turbulent in the central region between the two grids (the interaction region). For a detailed description of the experimental platform and the nature of the turbulence it generates, see Tzeferacos et al. (2018).

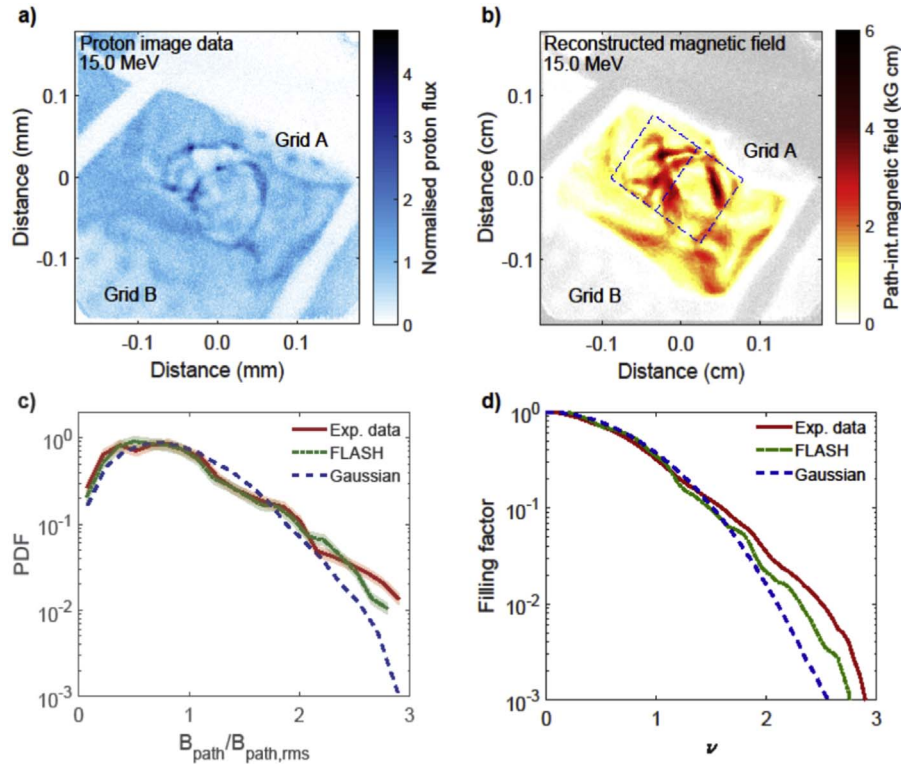
The electron density and temperature of the turbulent plasma are measured using collective Thomson scattering (Katz et al. 2012) and are found to be  $n_e \approx 9 \times 10^{19} \text{ cm}^{-3}$  and  $T_e \approx 400 \text{ eV}$  immediately after the formation of the turbulent region ( $t \approx 27 \text{ ns}$  after the start of the drive). The mean velocity ( $u_{\text{flow}}$ ) and the turbulent velocity ( $u_{\text{turb}}$ ) of the flow are also obtained by this diagnostic. Prior to collision, the counterpropagating flows reach velocities of  $u_{\text{flow}} \approx 2 \times 10^7 \text{ cm s}^{-1}$ , whereas in the turbulent region we measure at late times  $u_{\text{flow}} \approx 5 \times 10^6 \text{ cm s}^{-1}$  and  $u_{\text{turb}} \approx 10^7 \text{ cm s}^{-1}$  at the driving scale of the turbulence ( $L \approx 400 \mu\text{m}$ , set by the grid spacing).

The evolution of the plasma interaction region is determined using self-emitted soft X-rays (Figure 1(b) shows the plasma emission at  $38 \text{ ns}$  after the start of the laser drive). As discussed in Tzeferacos et al. (2018), fluctuations in the emissivity of such a plasma can be related to fluctuations of density (Churazov et al. 2012); the latter exhibit a Kolmogorov power-law spectrum, with a driving scale  $L$  consistent with the grid spacing detailed above (Figure 1(c)). The spatial extent of the interaction region over time can also be measured using the

X-ray diagnostic. Further details concerning the plasma state are given in Appendix B.

The stochastic magnetic fields amplified in the turbulent plasma (Tzeferacos et al. 2017, 2018) are measured using proton radiography (Figure 2). A  $420 \mu\text{m}$  diameter  $\text{SiO}_2$  capsule, with a  $2 \mu\text{m}$  thick shell, is filled with  $18 \text{ atm}$   $\text{D}^3\text{He}$  gas ( $6 \text{ atm } ^2\text{D}$  and  $12 \text{ atm } ^3\text{He}$ ) and is placed  $10 \text{ mm}$  away from the interaction region. The capsule is imploded using 17 beams (frequency-tripled to  $351 \text{ nm}$ , providing  $270 \text{ J/beam}$  for a  $1 \text{ ns}$  pulse) to produce  $3.3$  and  $15 \text{ MeV}$  fusion protons (Li et al. 2006; Kugland et al. 2012; Manuel et al. 2012). The protons are recorded on the opposite side of the capsule with a nuclear track detector (CR-39) film pack,  $27 \text{ cm}$  from the plasma interaction region, achieving a magnification of  $\times 28$ . Figure 2(a) shows a proton radiograph of the plasma corresponding to the same time as the X-ray image in Figure 1(b). The presence of strong inhomogeneities in the proton flux and the stochastic, nonregular morphology of the structures is due to protons being deflected by strong tangled magnetic fields. From the flux inhomogeneities observed in the proton image, the experimental radiographs can be inverted (Bott et al. 2017; Graziani et al. 2017) to recover two components of the path-integrated magnetic field (Figure 2(b)). Using the measured spatial extent of the interaction region, we can show that the measured path-integrated magnetic field corresponds to a root mean square (rms) value  $B_{\text{rms}} \approx 65\text{--}80 \text{ kG}$ , with a typical correlation length  $\ell_B \approx 90 \mu\text{m}$ . These experimental values are consistent with the results of the FLASH simulations, which give  $B_{\text{rms}} \approx 80\text{--}100 \text{ kG}$  and  $\ell_B \approx 50 \mu\text{m}$ , when the effects of diffusion of the imaging beam caused by small-scale magnetic fields and the underestimation of the magnetic energy by the reconstruction algorithm in the presence of small-scale caustics are taken into account (Tzeferacos et al. 2018).

The statistics of the path-integrated magnetic field are expected to deviate from Gaussian as a result of the spatial intermittency. To quantify this, we show the probability density



**Figure 2.** Magnetic field reconstruction. (a) 15 MeV proton radiography image of the entire interaction region at 38 ns, without a pinhole shield present in the path. For clarity, the image length scales are shown without the  $\times 28$  magnification factor; with this factor, the image has dimensions of  $10 \text{ cm} \times 10 \text{ cm}$ . (b) Magnitude of the two components of path-integrated magnetic field that are perpendicular to the proton-beam path, reconstructed (Bott et al. 2017) from the 15 MeV proton image in (a). (c) PDF of the magnitude of the path-integrated magnetic field  $B_{\text{path}}$  at 38 ns. The PDF (red) is calculated using the mean of the PDFs for the two rectangular regions depicted in (b); the uncertainty is derived from the standard error. The PDF of the path-integrated field arising in the FLASH simulations is also plotted (green), as is a Gaussian reference (blue) with the same rms field strength. (d) The fraction of area of the path-integrated magnetic field in which the field magnitude  $B_{\text{path}}$  satisfies  $B_{\text{path}} \geq \nu B_{\text{path,rms}}$ , where  $B_{\text{path,rms}}$  is the rms path-integrated field. This quantity is again calculated from the rectangular regions demarcated in (b).

function (PDF) of the magnitude of the path-integrated field in Figure 2(c) and the field filling factor in Figure 2(d). Both exhibit extended tails. Because the path-integrated magnetic field is spatially intermittent, it follows that the field itself must also be spatially intermittent. Indeed, the PDF of the magnetic field strength in the FLASH simulations exhibits an exponential tail. The simulations also demonstrate that the deviation from Gaussian statistics is more pronounced in the three-dimensional true fields than in the two-dimensional path-integrated field (see also Appendix B). Such non-Gaussian, spatially intermittent magnetic fields are expected to arise when the fluctuation dynamo is operating (Schekochihin et al. 2004).

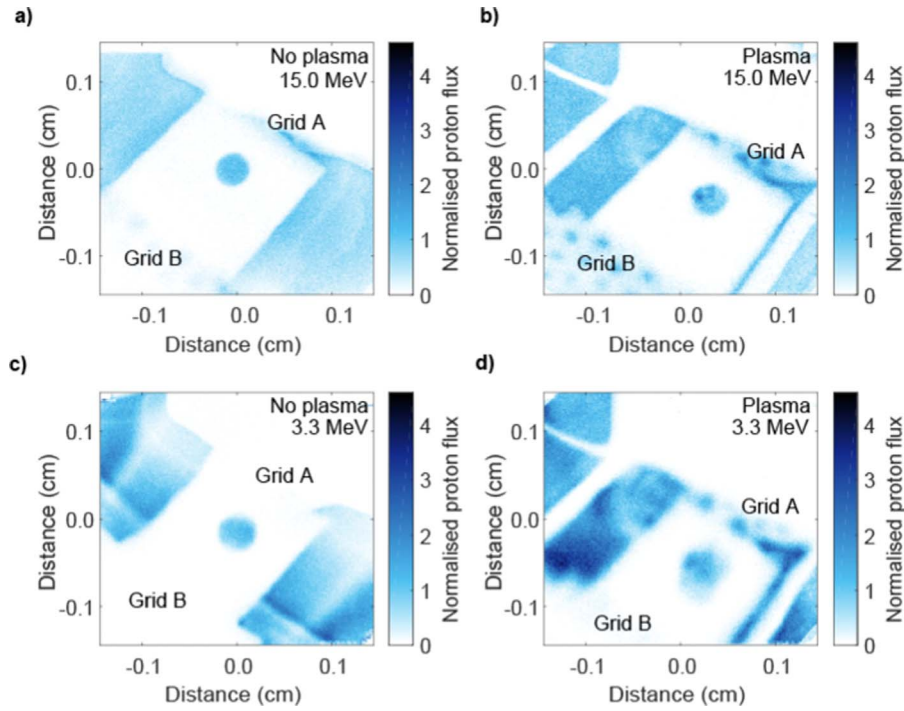
### 3. Transport Characterization

To characterize the transport of particles through the turbulent plasma, we modified our experimental platform to introduce a collimated proton beam. The collimation was achieved by placing a  $200 \mu\text{m}$  thick aluminum shield between the  $\text{D}^3\text{He}$  capsule and the interaction region, with a  $300 \mu\text{m}$  diameter pinhole (shown in Figure 1(a)). The pinhole imprint is then recorded on the detector plane, as shown in Figure 3. The proton-beam imprints appear deformed and broadened due to the interaction of the protons with the turbulent magnetized plasma. The proton-beam imprint contours are shown in Figures 4(a) and (b). The corresponding deflection velocities,  $\Delta v_{\perp}$ , as interpreted from the scattering angle  $|\Delta v_{\perp}|/V$ , where  $V$  is the proton-beam speed, are shown in Figure 4(c). Using the synthetic proton radiography diagnostic of the FLASH code

(Tzeferacos et al. 2017), we also post-processed the FLASH simulation results to recover proton trajectories and the resultant transverse deflections. These are in good agreement with the experimental measurements (Figure 4(c)).

The velocity deflection  $\Delta v_{\perp}$  due to magnetic fields scales independently of velocity, whereas the velocity deflection due to electric fields scales as  $\propto 1/V$  (shown in Appendix C). The near-equality of the deflection velocities of the two proton species, evident in Figure 4(c), suggests that scattering is predominantly due to magnetic fields. While there are many other possible processes that could lead to scattering of a charged-particle beam passing through a turbulent plasma, for our experiment we argue that these other processes are negligible, on account of the low density of the plasma and the high speed of the protons compared to driving-scale plasma motions. Detailed calculations and descriptions of possible electric field effects are given in Appendix D.

From our experimental measurement of  $\Delta v_{\perp}$ , we can calculate the associated scattering frequency in velocity-space,  $\nu \sim (\Delta v_{\perp}/V)^2/\tau$ , where  $\tau = \ell_i/V$  is the transit time of the particles through the plasma and  $\ell_i$  is the scale of the interaction region as inferred from the X-ray images. For a plasma with dimensions much greater than the proton mean free path  $\lambda \equiv V/\nu$ , our results imply an isotropic spatial diffusion coefficient  $\kappa \sim V^2/\nu = \ell_i V^3/(\Delta v_{\perp})^2$ . Because  $\kappa/V^3$  is constant in our experiment (Figure 4(d)), this implies  $(\Delta v_{\perp})^2 \propto \ell_i \propto \tau$ . This is consistent with a normal (Markov) spatial diffusion (Tsytovich 1977; Shalchi 2009; Subedi et al. 2017).



**Figure 3.** Proton pinhole images. Radiographs obtained on the CR-39 film pack with proton energies of (a) 15 MeV, with no plasma in the interaction region, and (b) 15 MeV, with a turbulent plasma in the interaction region. (c) Same as (a), but for the 3.3 MeV protons. (d) Same as (b), but for the 3.3 MeV protons. The pinhole shield is clearly seen to block most of the incoming proton flux from the capsule, and in the case where no plasma was present ((a) and (c)), it produces a fixed 300  $\mu\text{m}$  diameter beam of 3.3 and 15 MeV protons that passes through to the detector. For the case when a plasma is present in the interaction region ((b) and (d)), the beam is deformed and broadened before reaching the detector.

#### 4. Discussion

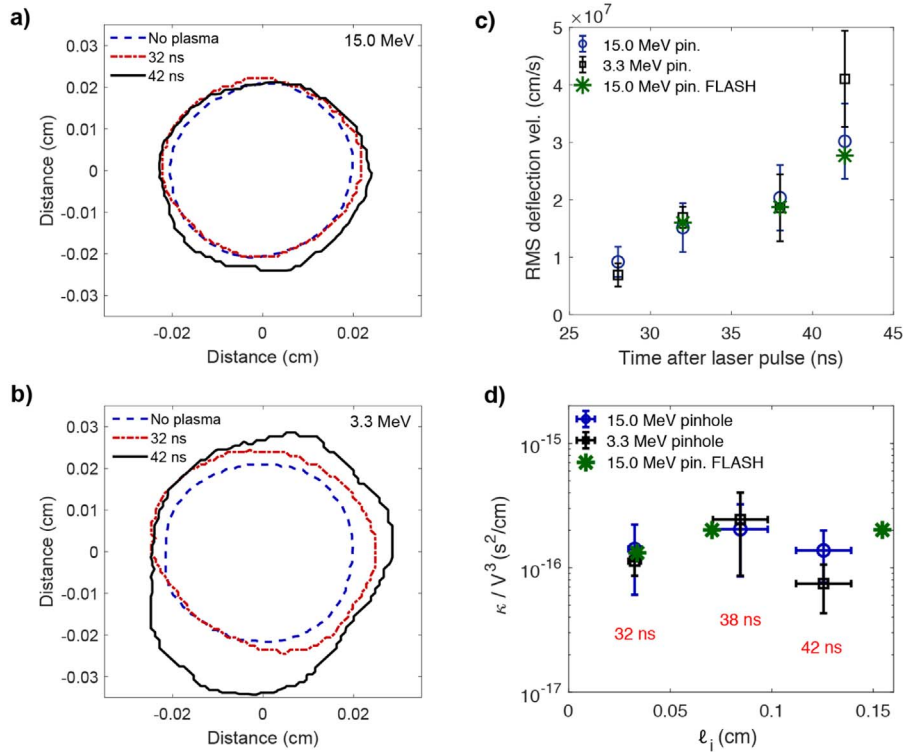
Because the charged-particle transport is consistent with normal spatial diffusion through a stochastic field, we can compare the experimental results to theoretical predictions from a random-walk process. Using characteristic values for the plasma properties corresponding to  $t = 38$  ns after the start of the drive, we take the size of the interaction region to be  $\ell_i \approx 0.08$  cm, the typical magnetic field strength  $B_{\text{rms}} \approx 100$  kG, and the correlation length  $\ell_B \approx 50$   $\mu\text{m}$ . For the case of normal diffusion, a random-walk argument gives  $\Delta v_{\perp} \approx q_e B_{\text{rms}} \sqrt{\ell_i \ell_B} / m_p \approx 1.9 \times 10^7$  cm s $^{-1}$  (see Appendix C), where  $m_p$  is the mass of a proton. This value is consistent with the measured rms deflection velocity (Figure 4(c)). Furthermore, because the values of  $V$ ,  $B_{\text{rms}}$ , and the power spectrum of the magnetic energy (and therefore the value of  $\ell_B$ ) do not change in the experiment after the magnetic field amplification saturates (see Tzeferacos et al. 2017, 2018 and Appendix A), the random-walk model also predicts a constant  $\kappa/V^3 \sim m_p^2 / (q_e B_{\text{rms}})^2 \ell_B \approx 1.9 \times 10^{-16}$  s $^2$  cm $^{-1}$ , in quantitative agreement with the experimental results (Figure 4(d)).

For isotropic statistics and  $r_g/\ell_B \gg 1$ , the proton mean free path is  $\lambda \approx 10^4$  cm. In this regime, theory (Dolginov & Toptygin 1967) and simulations (Subedi et al. 2017) predict that  $\lambda/\ell_B \propto (r_g/\ell_B)^2$ . The simulations of Subedi et al. (2017) predict a scaling coefficient of 1.5 and extend to  $r_g/\ell_B \approx 40$ . Extrapolating the results of Subedi et al. (2017) by a factor of 13 and 28 to the values  $r_g/\ell_B \approx 520$  for the 3.3 MeV protons and  $\approx 1, 100$  for the 15 MeV protons gives  $\lambda/\ell_B \approx 0.6 \times 10^6$  and  $\approx 2 \times 10^6$ . These values agree within a factor of order unity with the experimental value of  $\lambda/\ell_B \approx 2 \times 10^6$  that we obtain for the two proton energies. More importantly, our results demonstrate that, for the conditions present in the experiment (i.e., a beam with a diameter  $D > \ell_B$  of charged

particles in the  $r_g \gg \ell_B$  regime that traverses a stochastic and spatially intermittent magnetic field with a path length  $\ell_i > \ell_B$ ), the diffusion is not affected by the spatial intermittency of the stochastic magnetic fields. This is also demonstrated by the FLASH simulations of the experiment and numerical simulations presented in Appendix E.

The results of our experiments validate the use of standard diffusion theory in modeling the transport of UHECRs in the IGM, e.g., Kotera & Lemoine (2008), Globus et al. (2008, 2019), and Globus & Piran (2017), because all three of the above conditions are satisfied. This is useful in view of the increased interest in such modeling motivated by the recent detection by the Pierre Auger Observatory of a significant anisotropy in the arrival directions of cosmic rays of energy above 8 EeV (The Pierre Auger Collaboration 2017).

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**Figure 4.** Diffusive scattering of the proton beam. Contours of the beam imprint on the CR-39 plate for (a) 15 MeV protons and (b) 3.3 MeV protons, taken with different delay times after the start of the drive beams. (c) The rms transverse deflection velocity acquired by the proton beam, calculated using the contour analysis of the pinhole image (for both proton species), and evaluated for the FLASH simulations. (d) Measured spatial diffusion coefficient as a function of the plasma interaction length, as determined from the X-ray self-emission images and the 15 MeV pinhole synthetic radiographs from the FLASH simulations.

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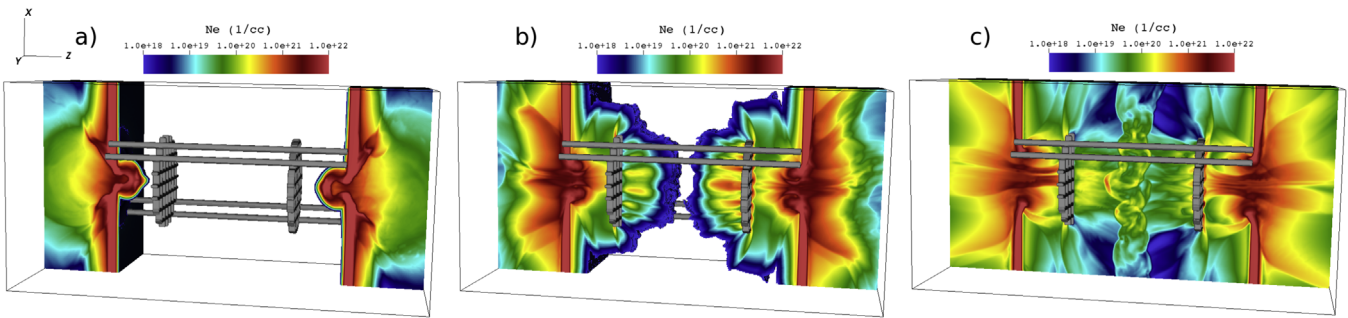
### Appendix A FLASH Simulations

We designed the experimental platform using three-dimensional radiation-magnetohydrodynamics (MHD) simulations carried out with the FLASH code. FLASH is a parallel, multiphysics, adaptive-mesh-refinement (AMR), finite-volume, high performance computing (HPC) Eulerian code (Fryxell et al. 2000; Dubey et al. 2009) that scales well to over a 100,000 processors. This is accomplished by exploiting a variety of in time parallelization techniques, and a combination of message passing and threading to optimally utilize hardware resources. The code is publicly available (<http://flash.uchicago.edu>) and has been successfully applied in a wide range of disciplines including astrophysics, cosmology, combustion, fluid dynamics, turbulence, and high-energy-density laboratory plasmas (HEDLP).

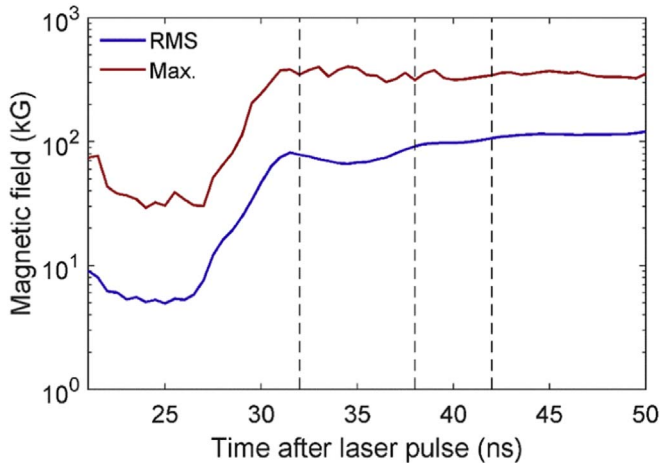
We performed a series of high-fidelity three-dimensional FLASH radiation-MHD simulations on the Mira supercomputer

at the Argonne National Laboratory. This simulation campaign to design the original experimental platform of Tzeferacos et al. (2018) is described in detail in Tzeferacos et al. (2017); considerations on the pinhole design are discussed in a companion paper. The simulations take advantage of the entire suite of HEDLP capabilities of FLASH (Tzeferacos et al. 2015), including its MHD solver (Lee 2013) extended to three temperatures (Tzeferacos et al. 2015), nonideal MHD effects such as magnetic resistivity (Tzeferacos et al. 2015), and Biermann battery (Fatenejad et al. 2013; Graziani et al. 2015), heat exchange between ions and electrons, implicit electron thermal conduction and radiation transport in the multigroup diffusion approximation, multitemperature tabulated equations of state and material opacities, and laser beams that are modeled using geometric-optics ray-tracing (Kaiser 2000) and deposit energy via inverse Bremsstrahlung.

We have used the FLASH code to design and simulate the experiment and, using the synthetic diagnostics routines of the code, constructed the proton radiographs shown in Figure C1. The latter are used to predict the values of  $\Delta v_{\perp}$  in Figure 4(c). The simulation domain spans 0.625 cm in the  $X$  and  $Y$  directions, and 1.25 cm along  $Z$ —the line of centers between the two targets. The spatial resolution is  $\sim 25 \mu\text{m}$ . The dynamics of the evolution are shown in Figure A1 and largely mirror the results discussed in Tzeferacos et al. (2017, 2018). The laser drive ablates the rear surfaces of the targets to launch two counterpropagating plasma flows (Figure A1(a)) that subsequently traverse the grids (Figure A1(b)), forming finger-like formations. The spatial offset of the grids by one grid aperture results in the two fronts interleaving to create a hot turbulent region in the center of the domain (Figure A1(c)).



**Figure A1.** Simulated temporal evolution. We show the logarithm of electron density (half-rendering), with gray contours denoting the grids and the supporting rods, at (a) 11 ns prior to collision (16 ns in the experiment), (b) 1 ns prior to collision (26 ns in the experiment), and (c) 16 ns after the collision (42 ns in the experiment) of the counterpropagating plasma flows.



**Figure A2.** Temporal evolution of  $B_{\max}$  and  $B_{\text{rms}}$ . Analysis of the FLASH simulations shows that the proton radiography images were taken (dashed vertical lines) after the saturation of the magnetic field amplification. Thus, the rms (blue solid line) and peak (red solid line) values of the magnetic field strength do not vary considerably in time.

Guided by the simulations, we modified the grids with respect to the design of Tzeferacos et al. (2018) and reduced the thickness of the wires from 300 to 100  $\mu\text{m}$ . This resulted in an increase in the number of apertures that in turn led to (i) increased throughput of kinetic energy and to (ii) the formation of a denser interaction region that is considerably thicker and more centered at the target chamber center than our previous design in Tzeferacos et al. (2017, 2018).

To quantify the temporal evolution of the magnetic field amplification in the simulations, we use a control volume (a cubic box of edge length 500  $\mu\text{m}$ , Tzeferacos et al. 2017) to initially track in time the propagating plasma front from grid A, and post-collision, the turbulent interaction region. In the control volume, we compute the peak and rms magnetic field strengths reported in Figure A2 with red and blue solid lines, respectively. The temporal trend agrees with the results of Tzeferacos et al. (2017, 2018). The time-series reveals that, after the amplification saturates, the magnetic field strength does not vary significantly in time and maintains values of  $B_{\text{rms}} \simeq 80\text{--}100$  kG and  $B_{\max} \simeq 350\text{--}400$  kG, with  $\ell_B \simeq 50$   $\mu\text{m}$ . This corroborates the arguments made in Section 3 regarding the scaling of the spatial diffusion coefficient  $\kappa$  because the proton images were taken after saturation (dashed vertical lines in Figure A2).

## Appendix B Plasma Characterization

In this section we describe the collective Thomson scattering diagnostic that was used to measure plasma properties, and report on measurements taken just after the collision of the two flows, and much later in time when the turbulent interaction region has developed. We then provide a comprehensive list of relevant plasma parameters, both measured and calculated, with particular care in accounting for the multi-ion nature of the plasma. We end this section by reporting on the evolution of the proton path length and the maximum path-integrated magnetic field strength as a function of time after the start of the drive laser pulse, which was key in the analysis described in Section 3.

### B.1. Thomson Scattering Diagnostic

A collective Thomson scattering (TS) diagnostic was employed to characterize the plasma properties. It was operated as either time-resolved or space-resolved (on different laser shots). For temporally resolved TS, the probe beam was a 1 ns pulse, delivering 25 J of 526.5 nm wavelength light to a 50  $\mu\text{m}$  cubed volume. For spatially resolved TS, the probe beam was a 0.6 ns pulse, delivering the same energy at the same wavelength to a cylindrical volume with a 50  $\mu\text{m}$  diameter cross-sectional area, and a 1.5 mm field of view (FOV). For both cases, the scattering angle was set to 63°, oriented with the scattering wavenumber parallel to the line of centers connecting the two targets and therefore the propagation axis of the two plasma flows. Figure B1 shows an example of the time-resolved measured data, and the best fit to the spectrum. The central peak at 526.5 ns corresponds to stray light from the probe beam. Fitting the scattering spectra and the total scattering power with the TS form factor (see Tzeferacos et al. 2018 for details) allows us to determine the values of the electron density ( $n_e$ ), the electron temperature ( $T_e$ ), the bulk flow velocity along the propagation axis of the two flows ( $u_{\text{flow}}$ ), and the turbulent velocity (i.e., fluctuations that broaden the scattering spectra) at the scale of the scattering volume ( $u_{\ell=50\mu\text{m}}$ ). While the turbulent velocity is measured from TS on scales  $\ell = 50$   $\mu\text{m}$ , we can find the turbulent velocity for the outer scales by assuming Kolmogorov scaling (as shown in Figure 2(d)) and see that  $u_{\text{turb}} \simeq (0.7\text{--}1.4) \times 10^7$   $\text{cm s}^{-1}$ . Before the collision (and until  $\sim 26$  ns), the speed of the flow of the bulk plasma is similar to the speed of the single jet, i.e.,  $u_{\text{jet}} \simeq u_{\text{flow}}$ , and naturally,  $u_{\text{flow}}$  decreases thereafter. These are given in Table 1 for 27 and 41 ns after the start of the laser pulse. The 27 ns data were taken using the

**Table 1**  
Thomson Scattering Data Plasma Properties Computed from Two Shots at 27 and 40 ns after the Laser Drive

Time after Drive	TS Mode	$T_e$	$n_e$	$u_{\text{flow}}$	$u_\ell = 50 \mu\text{m}$
27 ns	Spatially resolved	$400 \pm 80 \text{ eV}$	$9 \times 10^{19} \text{ cm}^{-3}$	$1.6 \times 10^7 \pm 2 \times 10^6 \text{ cm s}^{-1}$	...
41 ns	Temporally resolved	$330 \pm 50 \text{ eV}$	$1 \times 10^{20} \text{ cm}^{-3}$	$5 \times 10^6 \pm 2 \times 10^6 \text{ cm s}^{-1}$	$3.5\text{--}7 \times 10^6 \text{ cm s}^{-1}$

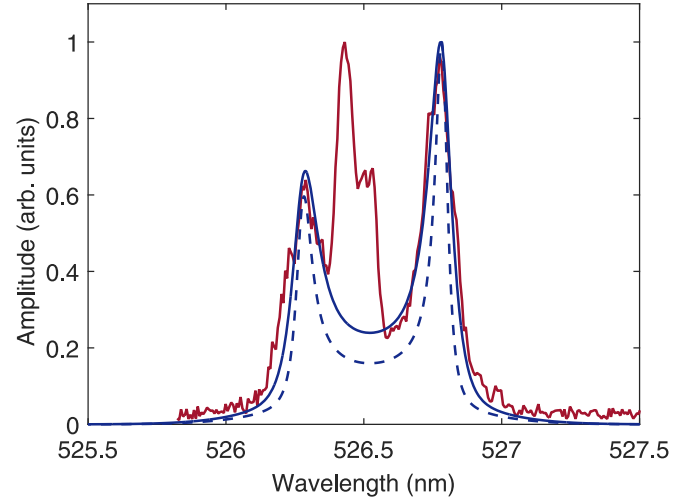
**Note.** The spatial and temporal variation for each value is reported here for the temperature and velocity measurements, and the electron density measurement has an error of 20%–30%.

**Table 2**  
Summary of Measured and Calculated Plasma Parameters

Plasma Parameter	Formula (S.I. units)	Value Time after Drive
Average atomic weight ( $\bar{m}$ )	C(50%), H(50%)	6.5 AMU
Temperature ( $T = T_e = T_i$ )		400 eV
Electron density ( $n_e$ )		$9 \times 10^{19} \text{ cm}^{-3}$
Unit charge ( $q$ )		1 Electron Charge
Carbon mass ( $m_C$ )		12 AMU
Carbon charge ( $Z_C$ )		6
Hydrogen charge ( $Z_H$ )		1
Average charge ( $\bar{Z}$ )	$\frac{1}{2}(Z_C + Z_H)$	3.5
Effective charge ( $Z_{\text{eff}}$ )	$\frac{1}{n_e}(Z_C^2 n_C + Z_H^2 n_H)$	5.3
Ion density ( $n_i$ )	$n_e / \bar{Z}$	$2 \times 10^{19} \text{ cm}^{-3}$
Respective ion species' densities ( $n_H, n_C$ )	$n_H = n_C = \frac{1}{2}n_i$	$1 \times 10^{19} \text{ cm}^{-3}$
Electron plasma frequency ( $\omega_{pe}$ )	$\sqrt{\frac{q^2 n_e}{\epsilon_0 m_e}}$	$5 \times 10^{14} \text{ s}^{-1}$
Outer scale ( $L$ )		400 $\mu\text{m}$
Jet velocity ( $u_{\text{jet}}$ )		$\sim 2 \times 10^7 \text{ cm s}^{-1}$
Turbulent velocity ( $u_{\text{turb}}$ )		$\sim 0.7\text{--}1.4 \times 10^7 \text{ cm s}^{-1}$
Carbon thermal velocity ( $v_{\text{th},C}$ )	$\sqrt{\frac{2T}{m_C}}$	
rms magnetic field ( $B_{\text{rms}}$ )		$\sim 80\text{--}100 \text{ kG}$
Debye Length ( $\lambda_D$ )	$\sqrt{\frac{T\epsilon_0}{n_e q^2} \left( \frac{1}{1 + Z_{\text{eff}}} \right)}$	$5 \times 10^{-6} \text{ cm}$
Electron specific heat ratio ( $\gamma_e$ )		1
Carbon-ion specific heat ratio ( $\gamma_C$ )		5/3
Carbon-ion sound speed ( $c_s$ )	$\sqrt{\frac{(Z_C k_B \gamma_C + \gamma_C) T}{m_C}}$	$1.55 \times 10^7 \text{ cm s}^{-1}$
Turbulence Mach number (Ma)	$u_{\text{turb}} / c_s$	$\sim 0.45\text{--}0.95$
Plasma $\beta$	$\frac{(n_e + n_i) T}{B_{\text{rms}}^2 / 2\mu_0}$	$\sim 160$
Carbon-carbon impact parameter ( $b_{\text{min},C}$ )	$\frac{Z_C^2 q^2}{8\pi\epsilon_0 T}$	$5 \times 10^{-9} \text{ cm}$
Electron-ion impact parameter ( $b_{\text{min},e}$ )	$\frac{\bar{Z} q^2}{8\pi\epsilon_0 T}$	$1 \times 10^{-9} \text{ cm}$
Coulomb logarithm carbon-carbon ( $\Lambda_{CC}$ )	$\log \left( \frac{\lambda_D}{b_{\text{min},C}} \right)$	6
Coulomb logarithm electron-ion ( $\Lambda_{ei}$ )	$\log \left( \frac{\lambda_D}{b_{\text{min},e}} \right)$	7
Carbon-carbon collision frequency ( $\nu_{CC}$ )	$4\pi \frac{Z_C^4 q^4}{(4\pi\epsilon_0)^2} \Lambda_{CC} \frac{n_C}{m_C^2 v_{\text{th},C}^3}$	$3 \times 10^{11} \text{ s}^{-1}$
Electron-electron collision time ( $\tau_{ee}$ )	$\frac{12\pi^{3/2}}{2^{1/2}} \frac{n_e q^4 \Lambda_{ei}}{m_e^{1/2} T_e^{3/2} \epsilon_0^2}$	$4 \times 10^{-12} \text{ s}$
Effective electron-ion collision frequency ( $\nu_{ei,\text{eff}}$ )	$Z_{\text{eff}} / \tau_{ee}$	$2 \times 10^{12} \text{ s}^{-1}$
Carbon gyrofrequency ( $\Omega_i$ )	$\frac{Z_C q B}{m_C}$	$5 \times 10^8 \text{ s}^{-1}$

**Table 2**  
(Continued)

Plasma Parameter	Formula (S.I. units)	Value Time after Drive
Electron gyrofrequency ( $\Omega_e$ )	$\frac{qB}{m_e}$	$2 \times 10^{12} \text{ s}^{-1}$
Dynamic viscosity ( $\mu$ )	$\frac{32\sqrt{2}}{15\pi^{3/2}} \frac{\sqrt{m\bar{H}}}{Z_C^2 (q/\sqrt{4\pi\epsilon_0})^4} \frac{T^{5/2}}{\Lambda_{ei}}$	$0.6 \text{ g cm}^{-1} \text{ s}^{-1}$
Kinematic viscosity ( $\nu_c$ )	$\mu / (n_i \bar{m})$	$2100 \text{ cm}^2 \text{ s}^{-1}$
Reynolds number (Re)	$\frac{u_{\text{turb}} L}{\nu}$	190
Viscous dissipation scale ( $\ell_{\nu}$ )	$\left( \frac{\nu^3}{\epsilon} \right)^{1/4} \sim \frac{L}{\text{Re}^{3/4}}$	8 $\mu\text{m}$
Spitzer conductivity ( $\sigma$ )	$4.1 \frac{n_e q^2}{m_e \nu_{ei,\text{eff}}}$	$87000 \text{ S cm}^{-1}$
Resistivity ( $\eta$ )	$\frac{1}{\sigma}$	$900 \text{ cm}^2 \text{ s}^{-1}$
Magnetic Reynolds number (Rm)	$\frac{u_{\text{turb}} L}{\eta}$	440
Magnetic Prandtl number (Pm)	$\frac{\text{Rm}}{\text{Re}}$	$\sim 1$
Resistive dissipation scale ( $\ell_\eta$ )	$\left( \frac{\eta^3}{\epsilon} \right)^{1/4} \sim \frac{\ell_\nu}{\text{Pm}^{1/2}}$	5 $\mu\text{m}$



**Figure B1.** Thomson scattering data. The measured ion acoustic spectrum at 41 ns after collision is shown in red, and the corresponding best fit is shown in solid blue. Table 1 gives values for the plasma parameters extracted from the fitting procedure. The dashed blue line is the best fit determined without inclusion of turbulent broadening.

spatially resolving configuration, and the errors in the table reflect the variation in the plasma parameters across the 1.5 mm extent of the collection cylinder. The electron density measurement has an error of about 20%–30%, primarily

determined by the uncertainty in the photometric calibration (Tzeferacos et al. 2018). It is evident that the average density and temperature in the interaction volume do not vary greatly over the duration of the experiment.

Plasma conditions derived using collective TS are complemented by proton radiography measurements, which provide maps of the magnetic field strength (Bott et al. 2017; Graziani et al. 2017). A comprehensive summary of measured and derived quantities needed to characterize the plasma state fully is given in Table 2. Viscosity and resistivity were calculated following the conventions in Dorf (2014) and Helender & Sigmar (2002), respectively. The viscosity depends on the effective electron-ion collision frequency, which accounts for the presence of multiple ion species present in the plasma. Viscous and resistive scales were determined according to the conventions given in Schekochihin & Cowley (2007). The sound speed was calculated for the carbon ion acoustic wave speed because this is lower than the hydrogen ion acoustic wave speed. From this, we see that the turbulence is subsonic.

We see that the plasma is collisional and all species are predominantly unmagnetized. We observe that for conditions occurring at the beginning of the collision between the two flows, the fluid Reynolds number (Re) and the magnetic Reynolds number (Rm) are of similar magnitude (i.e., the magnetic Prandtl number, Pm, is of order unity). As the flow evolves, the temperature decreases and Pm decreases as well. Therefore we expect to be in a low magnetic Prandtl-number regime,  $Pm \lesssim 1$ , in general.

### B.2. Time Evolution of Plasma

In Section 3 we measure the proton spatial diffusion coefficient through a stochastic magnetic field of a particular rms field strength. This measurement is taken at different times, corresponding to different extents of plasma expansion and therefore different path lengths. Thus we require a time-resolved measurement of the path length of the proton beam through the plasma,  $\ell_i$ . This measurement is shown in Figure B2(a).

We also need to know the rms field strength and correlation length of the stochastic magnetic field in the plasma; time-resolved measurements of these quantities are presented in Figure B2(b).

The measurement of the path length is obtained from the full width half-maximum of the plasma interaction observed in X-ray images. These are obtained by recording the plasma self-emission onto a framing camera with  $\sim 1$  ns gate width and filtered with  $0.5 \mu\text{m}$   $\text{C}_2\text{H}_4$  and  $0.15 \mu\text{m}$  Al. Spatial imaging is achieved using a  $50 \mu\text{m}$  diameter pinhole. Thus, the width of the plasma interaction region is defined as being the full width half-maximum of the X-ray emission profile in the direction parallel to the initial plasma flows. We then calculate the path length by projecting the distance onto the path of the proton beam through the interaction region (which has a known angle of  $55^\circ$  with respect to the direction of the initial plasma flows). The error in the measurement is determined from the standard deviation of the measured interaction-region width across the image. These measurements are confirmed by FLASH simulations (see above for a description of the FLASH code and the numerical setup), which show the width of the interaction

region increasing over time: the numerical values obtained directly from the simulation are shown in Figure B2(a), while Figure B3 illustrates the extent of the interaction region at times of 32, 38, and 42 ns.

We derive estimates of the rms field strength  $B_{\text{rms}}$  and the correlation length  $\ell_B$  of the stochastic magnetic field from the two spatially resolved components of the path-integrated magnetic field, which are themselves determined from full proton images of the turbulent plasma (see Figure 2); the analysis techniques used to perform these calculations are described in Bott et al. (2017). To assess the robustness of the result, we apply the technique to proton images at two additional times ( $t = 32, 42$  ns) after the initiation of the laser drive; these images (and the associated path-integrated fields) are shown in Figure B4. To derive an estimate of the uncertainty on the measurements, we perform these calculations for two rectangular regions, and use the means and errors to determine the results plotted in Figure B2(b). We find that within the experimental uncertainty of the measurement, the rms field strength and correlation length does not change in time. Furthermore, the experimental values are consistent with the FLASH simulation values that give  $B_{\text{rms}} \simeq 80\text{--}100$  kG and  $\ell_B \simeq 50 \mu\text{m}$ , when the effects of diffusion of the imaging beam caused by small-scale magnetic fields and the underestimation of the magnetic energy by the reconstruction algorithm in the presence of small-scale caustics are taken into account (Tzeferacos et al. 2018).

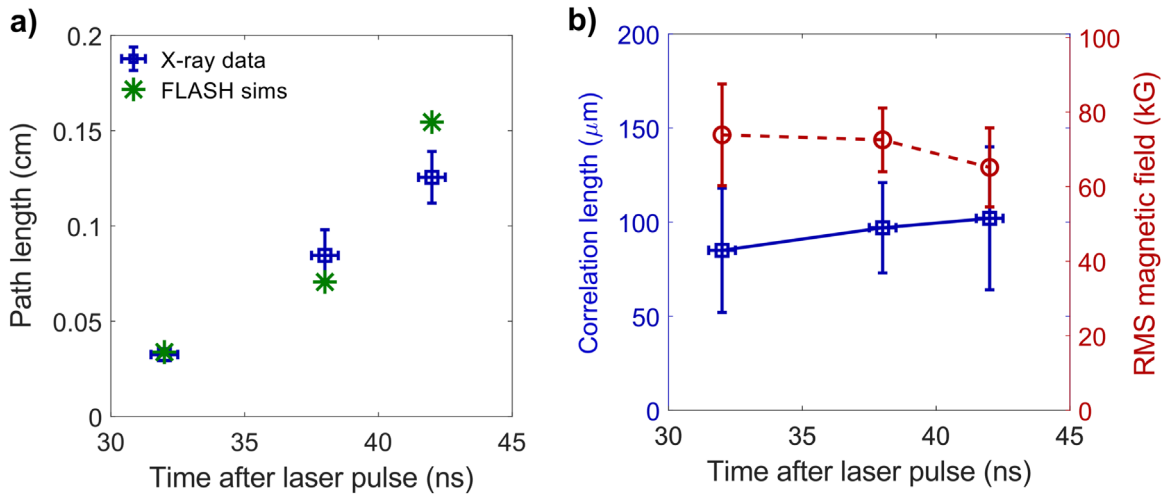
### B.3. Magnetic Field Spatial Intermittency

In Section 2 we stated that the observed spatial intermittency of the measured components of the path-integrated magnetic field provided evidence that the stochastic fields themselves were spatially intermittent; here, we demonstrate this claim explicitly for the FLASH simulations. Figures B5(a) and (b) show the PDFs and filling factors, respectively, of the true three-dimensional field, along with the equivalent quantities for a Gaussian stochastic field.

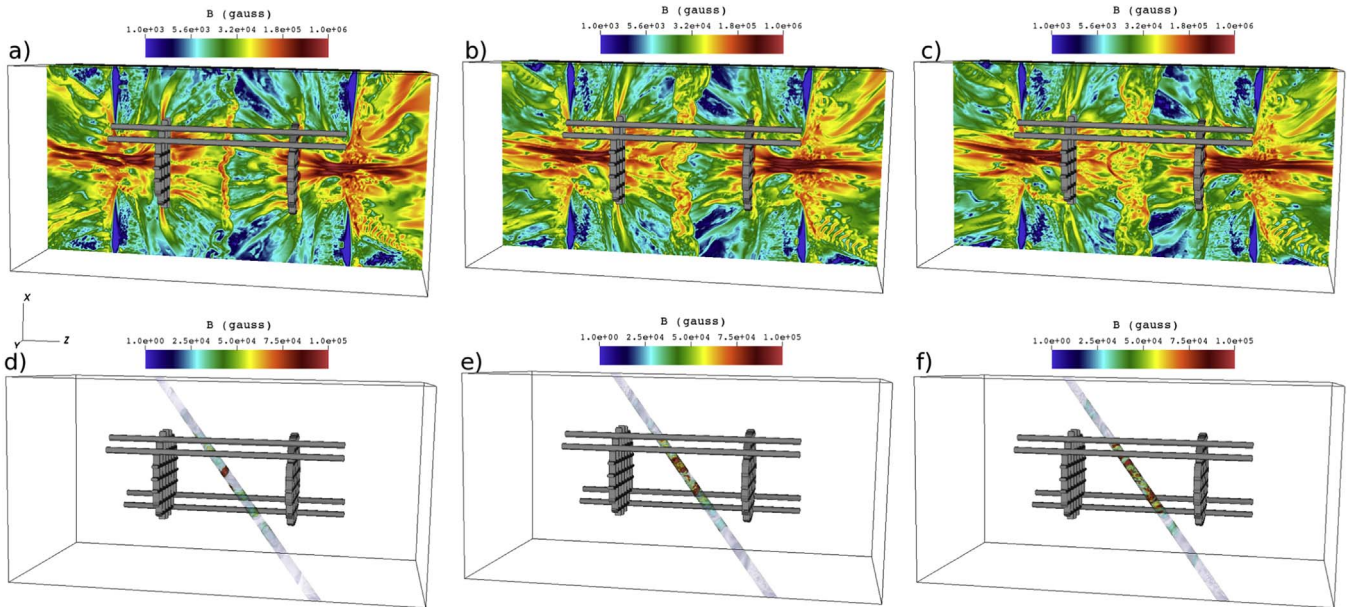
We observe that at all times in the simulation, the departure from Gaussian behavior is much more significant than the equivalent departure observed for the path-integrated field. In the simulations, for  $B/B_{\text{rms}} > 1$ , an exponential fit is obtained for the PDF; this matches the results of previous simulations of the turbulent dynamo (Schekochihin et al. 2004). To assess the robustness of our results, we also consider the PDF and filling factor of the experimentally recovered components of the path-integrated at 32 ns, and compare them to the equivalent quantities (at 38 ns) depicted in Section 2; the results are shown in Figure B6.

We find in both cases that there is a departure from Gaussian behavior at sufficiently high values of the path-integrated field magnitude  $B_{\text{path}}$  as compared to the rms value  $B_{\text{path,rms}}$ , and also for sufficiently high values of the filling factor  $\nu$ . However, we also observe some fluctuation in the precise shape of both the PDF and the filling factor, which in turn suggests that such fluctuations should not be regarded as being physically significant.

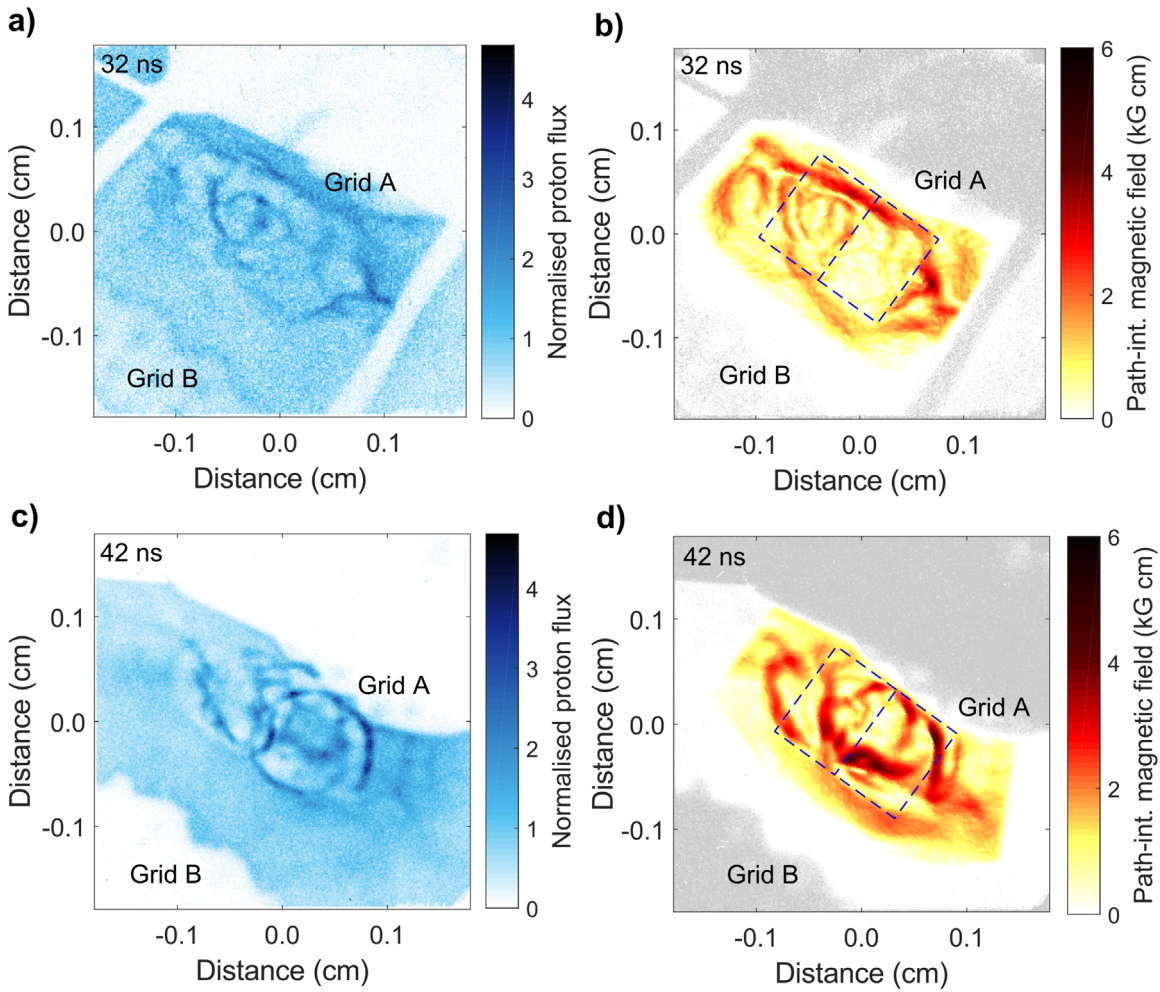




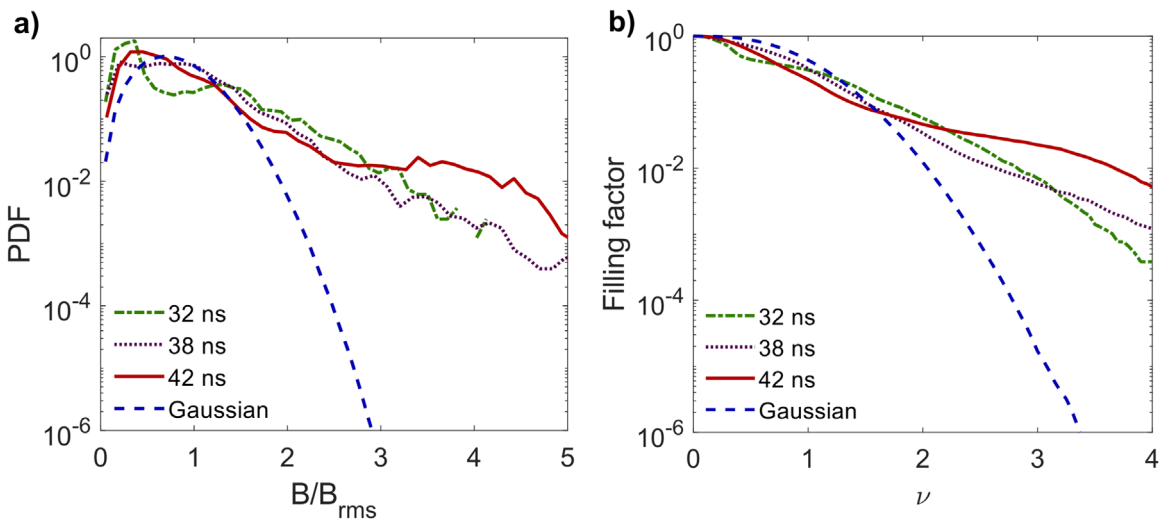
**Figure B2.** Evolution of the interaction region. (a) Plot of the proton-beam path length over time (shown in blue), as determined by measuring the full width half-maximum of the interaction region as seen in X-ray images. For comparison, the proton-beam path length as directly measured in the FLASH simulation is depicted by the green asterisks. (b) Plot of the correlation length (blue squares) and rms magnetic field strength (red circles) of stochastic magnetic fields present in the plasma over time, derived from the reconstructed path-integrated fields. The experimental values are consistent with the FLASH simulation values that give  $B_{\text{rms}} \simeq 80\text{--}100$  kG and  $\ell_B \simeq 50$   $\mu\text{m}$ , when the effects of diffusion of the imaging beam caused by small-scale magnetic fields and the underestimation of the magnetic energy by the reconstruction algorithm in the presence of small-scale caustics are taken into account (Tzeferacos et al. 2018).



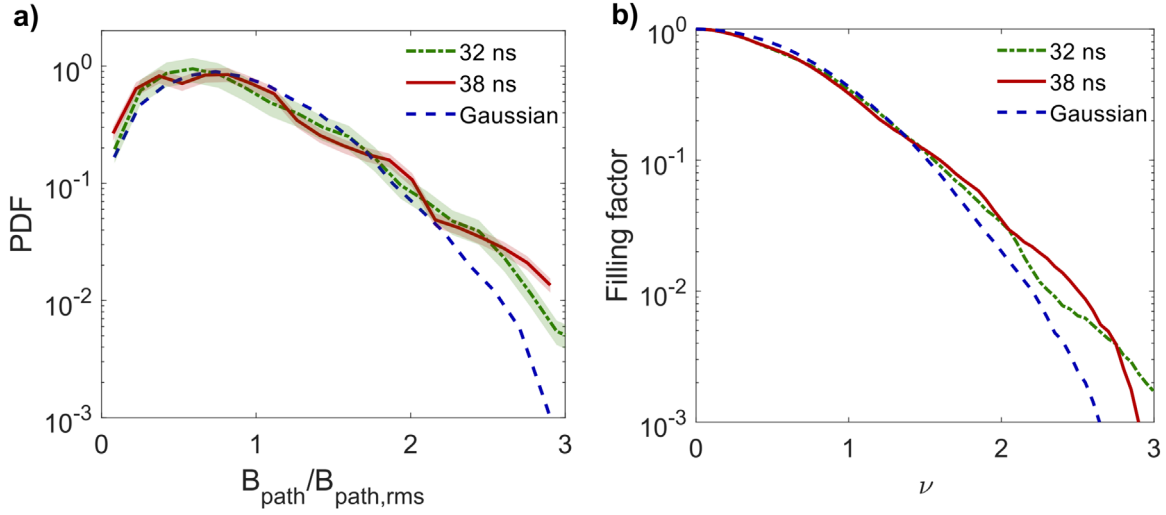
**Figure B3.** FLASH simulations of the magnetic field. Snapshots of the magnetic field evolution at times that correspond to (a) 32 ns, (b) 38 ns, and (c) 42 ns in the experiment. The intersections of the collimated proton beam with the magnetic field of the interaction region is shown for the same times in panels (d)–(f). We clearly see the expansion of the interaction region with time.



**Figure B4.** Full proton images of the turbulent plasma. (a) Proton image data for  $t = 32$  ns after the drive. (b) Magnitude of perpendicular path-integrated magnetic field at  $t = 32$  ns; the regions subsequently analyzed to determine the rms field strength and correlation length are demarcated in a blue dashed line. (c) Proton image data for  $t = 42$  ns after the drive. (d) Magnitude of the perpendicular path-integrated magnetic field at  $t = 42$  ns.



**Figure B5.** Spatial intermittency of FLASH-simulated magnetic fields. (a) Probability density function of the magnetic fields contained inside the sections of the cylindrical volumes depicted in Figures B3(d), (e), and (f), which are also contained within the full width half-maximum of the field with respect to the proton path length. A Gaussian with unit rms is plotted for reference. (b) The filling factor (see caption of Figure 2) for the same regions of the stochastic magnetic field.



**Figure B6.** Time evolution of path-integrated field spatial intermittency. (a) The PDF of the magnitude of the (perpendicular) path-integrated magnetic field measured using the full proton-imaging diagnostic at 32 and 38 ns. These quantities are calculated for the rectangular regions denoted in Figure B4(b) for  $t = 32$  ns and in Figure 2 for  $t = 38$  ns. As in Section 2, a reference PDF for a Gaussian path-integrated field is plotted. (b) The filling factor for the magnitude of the path-integrated field at 32 and 38 ns, along with the filling factor for a Gaussian random field.

### Appendix C Detailed Calculations of Proton Deflections

In this section we discuss the relation of the deflection velocity due to scattering angle,  $\Delta v_{\perp}$ , of the proton beam to the stochastic magnetic and electric fields present in the plasma. We also demonstrate through numerical tests that the excess blurring in the experimental proton image compared to the synthetic images from the reconstruction algorithm is due to diffusive scattering by small-scale magnetic field structures (see Figure 2). We then describe the procedures used to provide quantitative estimates of  $\Delta v_{\perp}$  shown in Figure 4(c) from both the magnetic field reconstruction technique paired with the smearing analysis of the driving-scale structures (illustrated in Figure 2) and from the pinhole-edge contour analysis from the pinhole-proton radiography (Figures 4(a) and (b)).

#### C.1. Calculation of Diffusive Deflection Velocities for Stochastic Electromagnetic Fields

The change in transverse velocity  $\Delta v_B$  due to angular scattering of a proton of initial speed  $V$  due to the presence of a magnetic field can be approximated in the small-deflections limit,  $|\Delta v_B| \ll V$ , by the following procedure: We begin by noting that the length of time that the particle experiences the magnetic field is given by  $\tau \approx z/V$ , where  $z$  is the path length. Making this substitution and integrating the proton equation of motion (with  $m_p$  the proton mass and  $q$  its charge), we find

$$\Delta v_B \approx \frac{q}{m_p} \hat{z} \times \int_0^{\ell_i} \mathbf{B}_{\perp} dz, \quad (\text{C1})$$

where  $\hat{z}$  is the unit vector of  $z$  and  $\mathbf{B}_{\perp}$  is the perpendicular component of the magnetic field relative to  $z$ . We see that  $\Delta v_B$  has no dependence on the initial velocity. The deflection angle is

$$\delta\theta_B = \frac{|\Delta v_B|}{V}, \quad (\text{2})$$

and the deflection length,  $\delta_B$ , which is the displacement of the proton from its original position due to deflection by the

magnetic field, is

$$\delta_B = r_{\text{det}} \delta\theta_B \propto \frac{1}{V}, \quad (\text{3})$$

where  $r_{\text{det}}$  is the distance from where this deflection occurred to the detector. Similarly, for electric fields, we have

$$\Delta v_E = \frac{q}{m_p V} \int_0^{\ell_i} \mathbf{E}_{\perp} dz \propto \frac{1}{V}, \quad (\text{4C})$$

and the deflection angle is

$$\delta\theta_E = \frac{|\Delta v_E|}{V}. \quad (\text{5C})$$

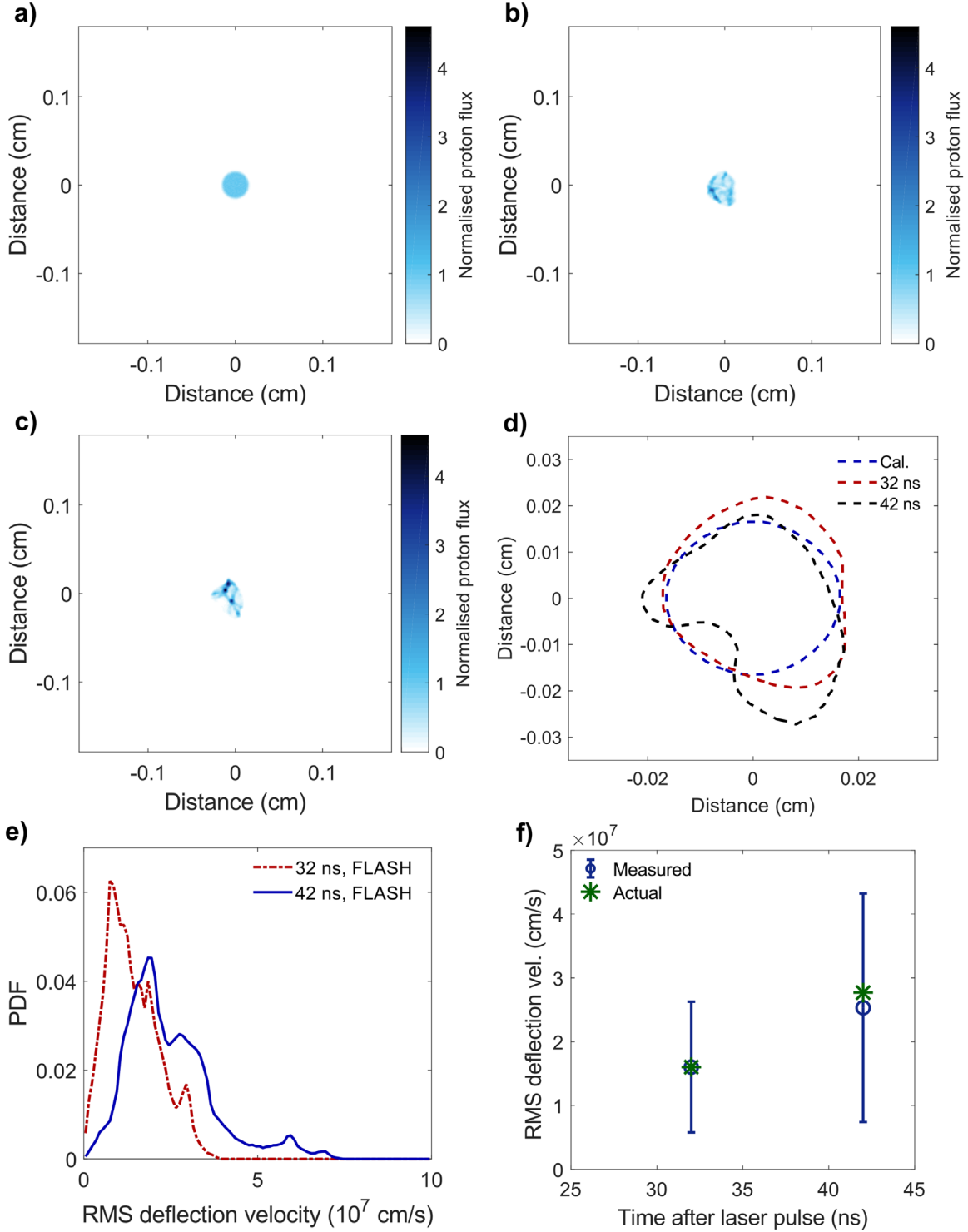
The deflection length due to electric fields is therefore

$$\delta_E = r_{\text{det}} \delta\theta_E \propto \frac{1}{V^2}. \quad (\text{C6})$$

After deriving  $\Delta v_{\perp}$  due to arbitrary magnetic and/or electric fields, we focus on stochastic fields. An expression for  $\Delta v_{\perp}$  as a function of the rms magnetic field strength  $B_{\text{rms}}$  and the field correlation scale  $\ell_B$  (this scale is the characteristic scale over which the magnetic field falls to zero from its peak value) is approximately half the size of the typical magnetic structure. The expression can be derived by the following heuristic argument. From Equation (C1), the change in velocity acquired by a proton traversing a single magnetic structure of size  $2\ell_B$  and strength  $B_{\text{rms}}$  is given by

$$\delta v_{\ell_B} \sim \frac{2qB_{\text{rms}}\ell_B}{m_p}. \quad (\text{7})$$

Assuming the velocity deflections add as a random walk, the overall velocity deflection can be approximated by  $\Delta v_B \simeq \delta v_{\ell_B} \sqrt{N_0}$ , where  $N_0 \gg 1$  is the number of structures encountered by the proton. Because  $N_0 \approx \ell_i/2 \ell_B$ , where  $\ell_i$  is the path length of the proton traversing the interaction region, it



**Figure C1.** Pinhole projections generated by FLASH simulations. Synthetic images of a collimated proton beam transiting through a turbulent plasma produced using FLASH. These correspond to (a) a calibration image with no plasma present and delay times of (b) 32 ns and (c) 42 ns. (d) The contours of these pinhole shapes were taken analogous to the contours in Figures 4(a) and (b). (e) Probability density function of a simulated proton beam passing through the FLASH simulation at 32 and 42 ns. (f) Estimate of rms deflection velocity as calculated from the FLASH pinhole-proton edge contours for 32 and 42 ns, as compared to the true values.

follows that  $\Delta v_B \sim \delta v_{\ell_B} \sqrt{\ell_i/2\ell_B}$ , leading to

$$\Delta v_B \approx \frac{qB_{\text{rms}} \sqrt{\ell_i \ell_B}}{m_p}, \quad (8)$$

where we have adjusted the numerical prefactor to agree with a rigorous correlation analysis (Bott et al. 2017). We note that  $\Delta v_B$  is independent of the initial proton velocity  $V$ , indicating that both proton species will experience the same diffusive deflection velocity due to scattering by the magnetic field. The

deflection angle due to the magnetic field is then

$$\delta\theta_B \approx \frac{\Delta v_B}{V} \approx \frac{qB_{\text{rms}}\sqrt{\ell_i\ell_B}}{m_p V}, \quad (\text{C9})$$

and the deflection length is

$$\delta_B \approx r_{\text{det}} \frac{qB_{\text{rms}}\sqrt{\ell_i\ell_B}}{m_p V}. \quad (\text{C10})$$

For a stochastic electric field with rms field strength of  $E_{\text{rms}}$  and scale of  $\ell_E$ , an expression similar to Equation (8) can be found,

$$\Delta v_E \approx \frac{qE_{\text{rms}}\sqrt{\ell_i\ell_E}}{m_p V}. \quad (\text{C11})$$

We note that because  $\Delta v_E \propto 1/V$ , faster protons will experience a lower diffusive deflection velocity due to electric field scattering.

The total deflection velocity due to scattering angle will have contributions from both magnetic and electric fields,

$$\Delta v_{\perp} = \Delta v_B + \Delta v_E. \quad (\text{C12})$$

For a single proton, it is impossible to distinguish between deflections due to magnetic and electric fields. However, in our experiment we generate protons of distinct energies (15.0 MeV protons from  $D^3\text{He}$  nuclear reactions, and 3.3 MeV protons from DD reactions). These protons experience the same fields; we now show that the relevant timescales allow this. The eddy turnover time at the driving scale is  $L/u_{\text{turb}} \simeq 4$  ns, while the shortest turnover time of plasma motions occurs at the dissipation scale  $\ell_{\eta}$ . Following the Kolmogorov scaling,  $u_{\text{turb},\ell_{\eta}} \sim u_{\text{turb}}(\ell_{\eta}/L)^{1/3}$ , which yields  $\ell_{\eta}/u_{\text{turb},\ell_{\eta}} \simeq 0.2$  ns, which is longer than the transit time of the proton beams across the plasma ( $\tau$ ), the pulse length of the beams ( $\tau_p$ ), and the time difference between the transit of the proton beams ( $\tau_d$ ). Specifically, for the 3.3 MeV DD protons,  $\tau \approx \ell_i/V_{\text{DD}} \approx 40$  ps,  $\tau_p \approx 150$  ps,  $\tau_d \approx r_i/V_{\text{DD}} - r_i/V_{D^3\text{He}} \approx 100$  ps ( $r_i$  is the distance from the proton source to the plasma). Thus, both proton beams sample the same fields, and so we can use the different velocity scalings illustrated by Equations (C1) and (C4) to distinguish between  $\Delta v_B$  and  $\Delta v_E$ .

### C.2. Experimental Measurement of Diffusive Deflection Velocity from Proton Pinhole-edge Contours

For the calculation of the pinhole-edge contours for both 15.0 and 3.3 MeV proton shown in Figures 4(a) and (b), we adopt a three-step procedure. First, the center of proton flux of the pinhole is calculated, using the flux distribution as a weight function. In principle, the presence of driving-scale inhomogeneities could mean that this is different to the projected center of the pinhole beam, given the position of the pinhole; however, for our data the difference between these points is negligibly small compared to the radius of the pinhole  $R = 150$   $\mu\text{m}$ . We then calculate the mean proton flux inside the specified pinhole radius,  $\Psi_0$  (taking the magnification into account). Finally, we define the contour of interest as the contour corresponding to 20% relative to the mean proton flux.

The motivation of this definition for pinhole-edge contour arises from two considerations: one theoretical, and one practical. Theoretically, deflections of the proton beam can be due to both driving-scale and small-scale structures. The former result in coherent distortions of the pinhole-edge contour from its undeflected position; the latter leads to a general expansion of the pinhole contour as a result of a shallowing of the gradient of the pinhole flux distribution at its edge. The 20% contour is sensitive to both effects (unlike the 50% contour, which is not sensitive to the small-scale structures). Practically, because of its large extent, the 20% contour is large enough to provide a reasonable average over the driving-scale structures. We do not choose an even lower percentage value for the contour because of distortions to such contours that arise from noise in the CR-39 plate in regions where the pinhole shield is present.

To extract quantitative estimates of  $\Delta v_{\perp}$ , the typical deflection velocity for a given proton-beam energy from the pinhole-edge contours, we compare the average deviation of the pinhole-edge contours obtained in the presence of the turbulent plasma with the calibration pinhole-edge contour for the same energy (i.e., with no plasma present). The comparison is performed numerically by measuring the squared displacement  $\Delta r^2(\theta) \equiv [r_{\text{plas}}(\theta) - r_{\text{cal}}(\theta)]^2$  between the plasma pinhole-edge contour radius  $r_{\text{plas}}(\theta)$  and the calibration plasma pinhole-edge contour radius  $r_{\text{cal}}(\theta)$  along a ray with polar angle  $\theta$  originating from the center of proton flux. The averaged displacement is then

$$\Delta r_{\text{rms}} \equiv \left[ \frac{1}{2\pi} \int_0^{2\pi} d\theta \Delta r^2(\theta) \right]^{1/2}. \quad (\text{C13})$$

The error in the displacement is in turn calculated from the variance of this measure. We find that small changes in the percentage value of the pinhole-edge contour do not significantly change our result.

We move from  $\Delta r_{\text{rms}}$  to a measurement of the typical deflection velocity using the fact that for any small-angle scattering process with typical deflection magnitude  $\Delta v_{\perp}$  and associated deflection length  $\delta \equiv r_{\text{det}} \Delta v_{\perp}/V$ , the effect of the diffusive scattering on the proton flux distribution can be modeled quantitatively by Bott et al. (2017)

$$\Psi(\mathbf{x}_{\perp}) = \int d^2\tilde{\mathbf{x}}_{\perp} \tilde{\Psi}(\tilde{\mathbf{x}}_{\perp}) \frac{1}{\pi\delta^2} \exp\left[-\left(\frac{\mathbf{x}_{\perp} - \tilde{\mathbf{x}}_{\perp}}{\delta}\right)^2\right], \quad (\text{C14})$$

where  $\mathbf{x}_{\perp}$  denotes perpendicular position on the detector,  $\tilde{\mathbf{x}}_{\perp}$  is a (two-dimensional) integration variable representing positions near  $\mathbf{x}_{\perp}$ ,  $\Psi(\mathbf{x}_{\perp})$  is the proton flux distribution at  $\mathbf{x}_{\perp}$ , and  $\tilde{\Psi}(\tilde{\mathbf{x}}_{\perp})$  is the proton flux distribution in the absence of the diffusive scattering process (evaluated at integration variable  $\tilde{\mathbf{x}}_{\perp}$ ). The model can be simplified under the assumption that the initial pinhole flux distribution is a uniform distribution, with radius  $R$  and mean flux  $\Psi_0$ . By symmetry, the smeared pinhole-proton flux distribution  $\Psi(\mathbf{x}_{\perp})$  is only a function of radial distance  $r$  from the pinhole center-of-mass under this assumption; that is,  $\Psi(\mathbf{x}_{\perp}) = \Psi(r)$ . Introducing polar coordinates ( $\tilde{r}$ ,  $\tilde{\theta}$ ) for integration variable  $\tilde{\mathbf{x}}_{\perp}$ , where  $\tilde{\theta}$  denotes the angle between vectors  $\mathbf{x}_{\perp}$

and  $\tilde{x}_\perp$ , Equation (C14) becomes

$$\Psi(r) = \frac{\Psi_0}{\pi\delta^2} \exp\left[-\left(\frac{r}{\delta}\right)^2\right] \int_0^R d\tilde{r} \tilde{r} \exp\left[-\left(\frac{\tilde{r}}{\delta}\right)^2\right] \times \int_{-\pi}^{\pi} d\tilde{\theta} \exp\left[\frac{2r\tilde{r}\cos\theta}{\delta^2}\right]. \quad (\text{C15})$$

Under the further assumption that the radius of the pinhole  $R$  is much greater than the smearing parameter  $\delta/\mathcal{M}$  (normalized by the magnification factor), we have that for values of the radial distance  $r \sim R$ , the dominant contribution to the integral in Equation (15) arises for  $\tilde{r} \sim r$ , and  $\theta \ll 1$ . These constraints give

$$\int_{-\pi}^{\pi} d\tilde{\theta} \exp\left[\frac{2r\tilde{r}\cos\theta}{\delta^2}\right] \approx \frac{\sqrt{\pi}\delta}{\sqrt{r\tilde{r}}}, \quad (\text{C16})$$

leaving

$$\Psi(r) \approx \frac{\Psi_0}{\sqrt{\pi}\delta} \int_0^R d\tilde{r} \sqrt{\frac{\tilde{r}}{r}} \exp\left[-\left(\frac{r-\tilde{r}}{\delta}\right)^2\right]. \quad (\text{C17})$$

Because the dominant contribution to the integral with respect to  $\tilde{r}$  comes from  $\tilde{r}$  satisfying  $\tilde{r} - r \sim \delta$ , we can approximate  $\tilde{r} \approx r$ ; Equation (C17) then simplifies to

$$\begin{aligned} \Psi(r) &\approx \frac{\Psi_0}{\sqrt{\pi}\delta} \int_0^R d\tilde{r} \exp\left[-\left(\frac{r-\tilde{r}}{\delta}\right)^2\right] \\ &\approx \frac{\Psi_0}{2} \operatorname{erfc}\left(\frac{r-R}{\delta}\right), \end{aligned} \quad (\text{C18})$$

for  $\operatorname{erfc}(x)$  the complementary error function. We conclude that under this model, the 20% contour on average will satisfy

$$\frac{\Psi_0}{5} \approx \frac{\Psi_0}{2} \operatorname{erfc}\left(\frac{\Delta r_{\text{rms}}}{\delta_{\text{rms}}}\right), \quad (\text{C19})$$

where  $\delta_{\text{rms}}$  is the rms deflection length. It then follows that

$$\delta_{\text{rms}} \approx \Delta r_{\text{rms}} / \operatorname{erfc}^{-1}\left(\frac{2}{5}\right) \approx 1.65 \Delta r_{\text{rms}}. \quad (\text{C20})$$

### C.3. FLASH Simulations of Proton Diffusion

FLASH simulations are used to create synthetic images of the proton flux through a 300  $\mu\text{m}$  diameter pinhole, recorded on the CR-39 plate. This is shown in Figures C1(a)–(c) for different probing times. The collimated beam consists of 500,000 simulated protons. The resolution of the FLASH synthetic radiographs, determined by the imposed capsule smearing and the proton binning, is 50  $\mu\text{m}$  to match the experimental resolution. Next, we evaluate the 20% pinhole-edge contour in the same manner we have applied to the experimental data, and show the results in Figure C1(d). Notably, the profile at 42 ns after the laser drive exhibits significant anisotropy. This is naturally expected—in probing a relatively small area in a stochastic environment, it is reasonable that we would at some point probe an anisotropic portion in the magnetic field distribution. In fact, in the experimental data, shown in Figure 4(b), we can also see anisotropies in the flux distribution at 42 ns.

In the simulation, we can evaluate  $\Delta v_\perp$  directly for each proton by measuring its spatial displacement on the recording

screen with respect to the point the proton would have landed with no plasma present, correcting for the spatial shift of the center of the pinhole-proton flux. This allows us to determine the rms  $\Delta v_\perp$  as experienced collectively by all protons passing through the pinhole. We see in Figure 4(c) that FLASH predicts values of  $\Delta v_\perp$  in very close agreement with the experimental ones. This result further corroborates our previous explanation that the observed enhanced scattering is the result of a longer path-length through the stochastic magnetic fields in the turbulent plasma because the simulated proton beam is only subject to Lorentz forces and not to other processes (such as target charging or collisional broadening). In addition to the measurement of the rms  $\Delta v_\perp$ , for the simulation we are able to plot the PDF of deflection velocities; the result is shown in Figure C1(e). As expected, we see a peak around the rms value, in addition to a tail of higher values.

Finally, for the simulated data, we can apply the same analysis technique described and applied to the experimental data in the previous section to the synthetic pinhole-proton edge contours derived from FLASH. The results are shown in Figure C1(f). The estimates obtained are comparable to the directly measured values. Because of the anisotropies present in the simulated pinhole-proton contours, the uncertainty on the technique (calculated in the same way) is more significant than for the experimental data.

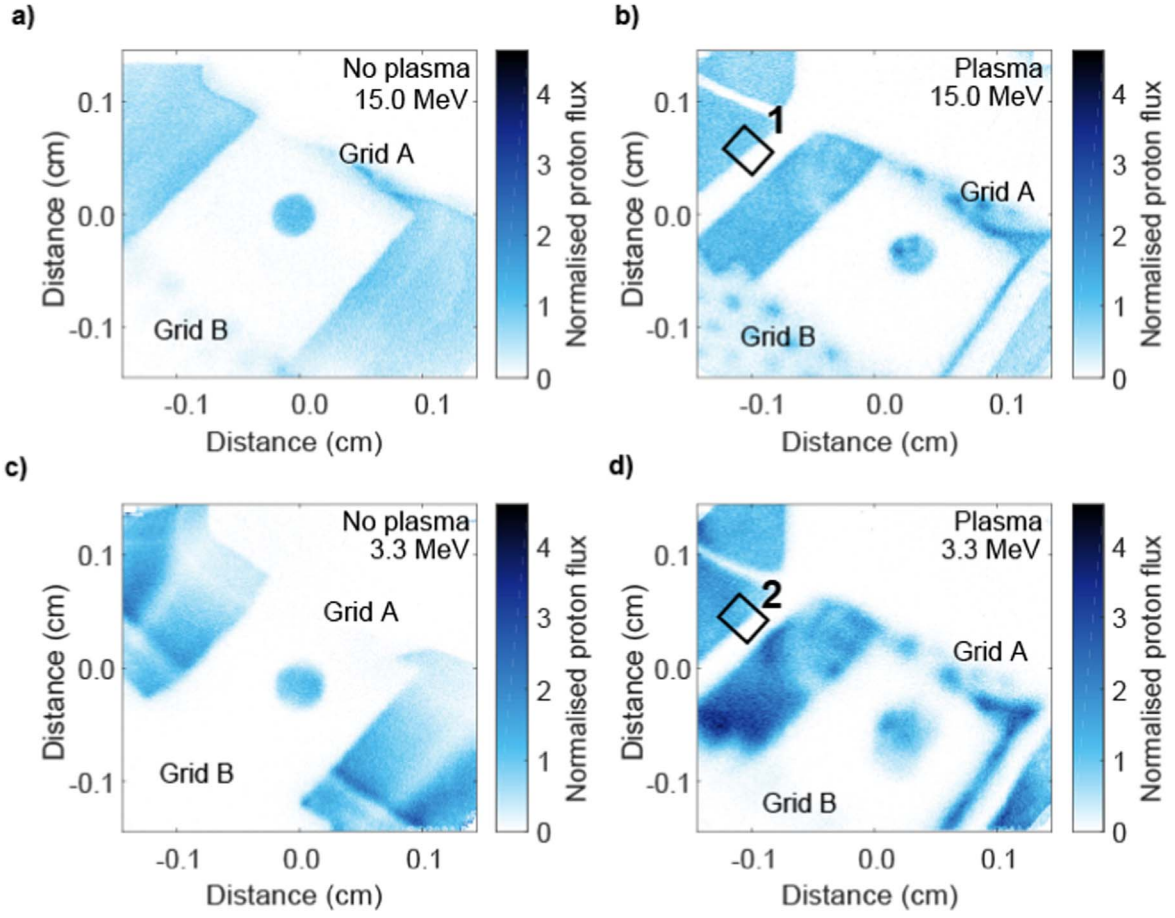
## Appendix D Physical Processes Resulting in Smearing of the Proton Beam

We find that most likely, the diffusive scattering is a result of stochastic magnetic fields. Assuming classical velocity-space diffusion, the estimate for the scattering velocity is obtained via a random-walk approximation, i.e., the sum of velocity deflections  $\delta v_{\ell_B} \approx \omega_g \ell_B$ , where  $\omega_g \equiv q_e B_{\text{rms}}/m_p$  is the gyrofrequency of the protons,  $B_{\text{rms}} \equiv \sqrt{\langle B^2 \rangle}$  is the measured rms magnetic field,  $q_e$  is the electron charge, and  $m_p$  is the proton mass; this results in  $\Delta v_\perp \approx \omega_g \sqrt{\ell_i \ell_B} \approx q_e B_{\text{rms}} \sqrt{\ell_i \ell_B}/m_p$ , where  $\ell_i$  is the size of the interaction region. As mentioned in the text, when we take  $\ell_i \sim 0.8$  mm,  $B_{\text{rms}} \sim 100$  kG and  $\ell_B \sim 50$   $\mu\text{m}$ , we find  $\Delta v_\perp \sim 1.9 \times 10^7$  cm s<sup>-1</sup>, which is indeed consistent with the observed rms deflection velocity in Figure 4(c).

As mentioned in Section 2, there are processes in addition to stochastic magnetic or electric fields that could result in the smearing of features on the proton image. In the remaining section, we discuss estimates of these effects in more detail and show that they are negligible.

### D.1. Capsule Size and Proton Velocity Uncertainty

The proton source, originating from an imploded capsule containing D<sup>3</sup>He and D<sub>2</sub> gas, has a finite source size, as well as a spread in emitted proton energies. The effective source size  $d_c$  of the capsule for 15.0 MeV protons is 45  $\mu\text{m}$  (Li et al. 2006). This is smaller than the initial capsule size because fusion reaction only occurs in the hot region that is formed well inside the imploding capsule. Including the magnification,  $\mathcal{M}$ , of the proton-imaging diagnostic, we conclude that all structures with sizes below  $\mathcal{M}d_c$  will be smeared out. This is consistent with the slope at the edge of the pinhole image for the calibration shots seen in Figure 4(a). The effect source size for 3.3 MeV protons is believed to be somewhat larger, an observation born



**Figure D1.** Charging of the target setup. Radially averaged lineouts from the flux-weighted center of the pinhole in the case of no plasma present in the interaction region are shown for (a) 15.0 MeV and (c) 3.3 MeV protons. This is compared to lineouts taken across an edge section of the target package from a proton radiograph in which a plasma was present in the interaction region, demarcated by 1 and 2 for (b) 15.0 MeV and (d) 3.3 MeV protons.

out in Figure 4(b): the 3.3 MeV calibration profile is broader than the 15.0 MeV one.

The blur resulting from uncertainties in the initial proton energy can be estimated as

$$\delta_v = r_{\text{det}} \delta \theta_B \frac{\delta V}{V} \approx 40 \mu\text{m}, \quad (\text{D1})$$

and as  $\delta_B \gg 40 \mu\text{m}$ , this is negligible. The term  $\delta V/V$  is calculated from the experimentally determined energy uncertainties  $\delta W/W$ : these are  $\sim 3\%$  and  $\sim 4.5\%$  for  $\text{D}^3\text{He}$  and  $\text{DD}$  protons, respectively. When we use  $\delta W/W \approx 2\delta V/V$ , this gives estimates for  $\delta V/V$  as  $\sim 1.5\%$  for 15.0 MeV protons and  $\sim 2.25\%$  for 3.3 MeV. Thus, this second effect is likely to be negligibly small.

In short, given the relatively small size of diffusive scattering due to stochastic magnetic fields compared to the finite source size, accounting for the source size is essential. As described in the previous section, we carry out our analysis using experimental pinhole images created in the absence of a plasma, and so we automatically account for these effects.

#### D.2. Electric Fields Due to Charging of the Target Package

In principle, it is possible that charging of the target package (e.g., due to high-energy electrons produced by laser-plasma processes) could allow for the emergence of an electric field. To rule this effect out, we compare a flux lineout across an edge

section of the target package from a proton radiograph in which a plasma was present in the interaction region (shown in Figures D1(b) and (d), labeled 1 and 2 for each proton species) to a radially averaged lineout from the flux-weighted center of the pinhole for when no plasma is present in the interaction region (shown in Figures D1(a) and (c)). These lineouts, shown in Figure D2, have comparable slopes, indicating any electric field due to charging is negligible in our analysis.

#### D.3. Electric Fields Generated by Plasma Fluid Motions

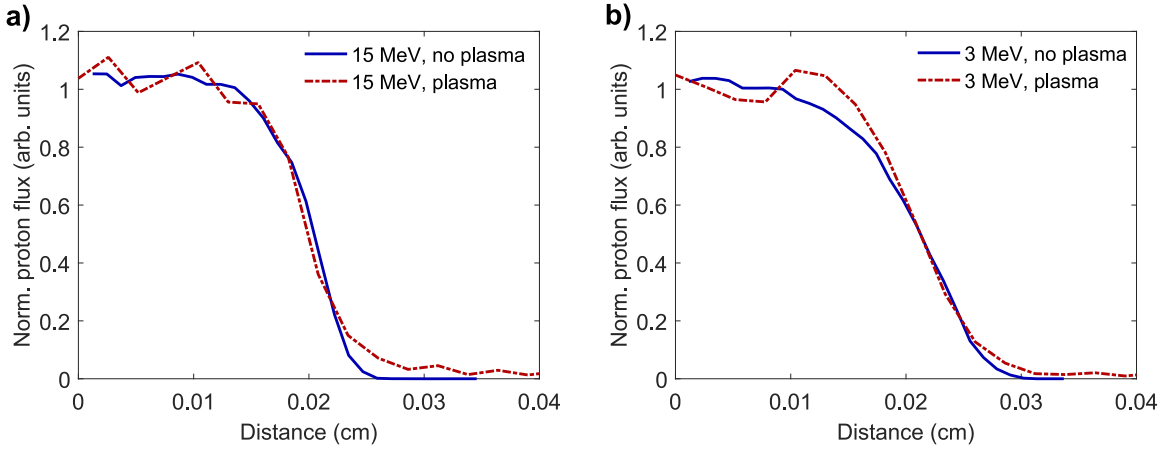
In Section 3 we asserted that scattering due to stochastic electric fields generated by turbulent motions of the electric field are small. This estimate is made by combining Equation (11C) with an appropriate estimate for  $E_{\text{rms}} \sim u_{\text{jet}} B_{\text{rms}}$ , and  $\ell_E \sim \ell_B$ . Then, it follows that

$$\Delta v_E \sim 1.6 \times 10^5 \left( \frac{u_{\text{jet}}}{2 \times 10^7 \text{ cm s}^{-1}} \right) \left( \frac{B_{\text{rms}}}{100 \text{ kG}} \right) \times \left( \frac{V}{2.5 \times 10^9 \text{ cm s}^{-1}} \right)^{-1} \left[ \frac{\ell_i \ell_B}{4 \times 10^{-4} \text{ cm}^2} \right]^{1/2} \text{ cm s}^{-1}, \quad (\text{D2})$$

as stated in Section 2.

#### D.4. Transverse Diffusion via Collisions

As the plasma is neither cold nor in thermal equilibrium, it is not straightforward to do an analytical calculation of transverse



**Figure D2.** Charging effects on pinhole. Lineouts across an edge section of the target package from a proton radiograph in which there was a plasma present in the interaction region (red, dashed) vs. a radially averaged lineout of the pinhole in the case of no plasma present in the interaction region (blue, solid) for (a) 15.0 MeV and (b) 3.3 MeV protons. It is clear that the slopes of the two cases are so similar that an electric field due to charging of the target is negligible.

diffusion due to collisional scattering. Under the classical approach of setting a background particle's Coulomb potential equal to the kinetic energy of an incoming test particle, we obtain estimates that can vary by a factor of order unity (Trubnikov 1965). A quantum-mechanical treatment is required that also takes into account long-range interactions, and we therefore used SRIM (Ziegler 1999) to simulate the perpendicular deflection of 3.3 and 15 MeV protons through 3.7 mm (the largest reported  $\ell_i$  at very late times) of a CH plasma of density  $9 \times 10^{19}$ . We found that this deflection corresponds to  $\Delta v_{\perp} = 2 \times 10^6$  cm s $^{-1}$  for DD protons and  $\Delta v_{\perp} = 8 \times 10^5$  cm s $^{-1}$  and D $^3$ He protons. Again, this is an order of magnitude lower than the measured  $\Delta v_{\perp}$  and is therefore negligible in our analysis.

#### D.5. Electrostatic Beam Instabilities

It is well known that in general, a beam of protons can drive kinetic instabilities. To determine whether this could occur for our experiment, we outline here an analysis investigating the kinetic stability of the combined plasma species and beam distribution functions. Linear kinetic stability for electrostatic modes can be established by considering the zeros of the dielectric function given by Boyd & Sanderson (2003) and Krall & Trivelpiece (1973),

$$\epsilon(p, \mathbf{k}) = 1 - \frac{\omega_{pe}^2}{k^2} \int_{C_L} dv_z \frac{d\bar{F}(v_z)/dv_z}{v_z - ip/k}, \quad (\text{D3})$$

where  $p$  is complex frequency,  $\omega_{pe}$  is the electron plasma frequency,  $v_z$  is the velocity of a particle in the one-dimensional distribution function,  $C_L$  is the contour of interest in momentum space, and the weighted distribution function,  $\bar{F}$ , is defined as

$$\bar{F} \equiv \frac{1}{n_e} \sum_{\alpha} Z_{\alpha}^2 \frac{m_e}{m_{\alpha}} F_{\alpha} \quad (\text{D4})$$

for mass,  $m$ , atomic charge,  $Z$ , and distribution function,  $F$  for all respective species, denoted by  $\alpha$ . It can then be shown that any isotropic monotonically decreasing distribution function is inherently stable against infinitesimal (linear) electrostatic perturbations.

The plasma in our experiment is collisional and contains two ion species (carbon and hydrogen) in equal ratio. We therefore assume that electron, hydrogen, and carbon have the following respective Maxwellian distributions:

$$F_e(v_z) = \frac{n_e}{\pi^{1/2} v_{th,e}} \exp[-(v_z/v_{th,e})^2], \quad (\text{D5})$$

$$F_H(v_z) = \frac{n_H}{\pi^{1/2} v_{th,H}} \exp[-(v_z/v_{th,H})^2] + \frac{n_b}{\pi^{1/2} \delta V} \exp[-((v_z - V)/\delta V)^2], \quad (\text{D6})$$

$$F_C(v_z) = \frac{n_C}{\pi^{1/2} v_{th,C}} \exp[-(v_z/v_{th,C})^2], \quad (\text{D7})$$

where  $v_{th,e} = \sqrt{2T/m_e}$  is the thermal electron velocity,  $v_{th,H} = \sqrt{2T/m_p}$  is the thermal plasma proton velocity,  $v_{th,C} = \sqrt{2T/m_C}$  is the plasma carbon ion velocity,  $V$  is the beam velocity,  $\delta V$  is the (thermal) width of the beam,  $n_e$  is the electron density,  $n_H$  is the plasma proton density,  $n_b$  is the beam proton density, and  $n_C$  is the plasma carbon ion density. We have taken the ion and electron temperatures to be the same:  $T_e = T_i = T$  (Tzeferacos et al. 2018). Assuming quasineutrality, we have

$$n_e = n_H + 6 n_C. \quad (\text{D8})$$

We have measured (see Appendix B on TS diagnostics)  $n_e = 9 \times 10^{19}$  cm $^{-3}$ , so  $n_H = n_C = n_e/7 = 1.28 \times 10^{19}$  cm $^{-3}$ . The total proton count recorded on the CR-39 detector is  $\sim 25,000$  for 15.0 MeV protons, and  $\sim 50,000$  for 3.3 MeV protons. Because the proton pulse duration is 150 ps, this gives a particle density in the proton beam of  $n_b \approx 10^8$  cm $^{-3}$ .

We now take the weighted distribution using Equation (D4) and examine  $\frac{d\bar{F}}{dv_z}$  to see if a bump in  $\bar{F}$  exists, i.e., if  $\frac{d\bar{F}}{dv_z}$  becomes positive anywhere. If it does not, then the distribution function



is monotonically decreasing. We have

$$\begin{aligned} \frac{d\bar{F}}{dv_z} &= \frac{1}{n_e} \frac{n_e}{\pi^{1/2} v_{th,e}} \frac{-2v_z}{v_{th,e}^2} \exp[-(v_z/v_{th,e})^2] \\ &+ \frac{1}{n_e} \frac{m_e}{m_p} \frac{n_H}{\pi^{1/2} v_{th,H}} \frac{-2v_z}{v_{th,H}^2} \exp[-(v_z/v_{th,H})^2] \\ &+ \frac{1}{n_e} \frac{m_e}{m_p} \frac{n_b}{\pi^{1/2} \delta V} \frac{-2(v_z - V)}{\delta V^2} \exp[-((v_z - V)/\delta V)^2] \\ &+ \frac{1}{n_e} \frac{m_e}{m_C} \frac{n_C}{\pi^{1/2} v_{th,C}} \frac{-2v_z}{v_{th,C}^2} \exp[-(v_z/v_{th,C})^2]. \end{aligned}$$

The distribution is clearly dominated by the electron distribution function. Any bump in the weighted distribution function must occur near the maximum of the beam distribution function—that is, when  $v_z \sim V$ . For such  $v_z$ , the contributions to  $\frac{d\bar{F}}{dv_z}$  from the plasma proton and carbon species is exponentially small ( $V \gg v_{th,H}, v_{th,C}$ ) and therefore negligible. Therefore, we can approximate  $\frac{d\bar{F}}{dv_z}$  for  $v_z \sim V$  as

$$\begin{aligned} \frac{d\bar{F}}{dv_z} &\approx \frac{-2v_z}{v_{th,e}^3} \exp[-(v_z/v_{th,e})^2] \\ &+ \frac{n_b}{n_e} \frac{m_e}{m_p} \frac{2(V - v_z)}{\delta V^3} \exp[-((v_z - V)/\delta V)^2]. \end{aligned} \quad (D9)$$

The global maximum of the second term on the right occurs when  $v_z = V - v_{th,b}/\sqrt{2}$ , where  $v_{th,b}$  is the beam thermal velocity (given by a spread in energies of about  $\sim 280$  eV), and can be recast as

$$\begin{aligned} &\frac{n_b}{n_e} \frac{m_e}{m_p} \frac{2(V - v_z)}{\delta V^3} \exp[-((v_z - V)/v_{th,b})^2] \\ &= \frac{n_b}{n_e} \frac{m_e}{m_p} \frac{\sqrt{2}}{\delta V^2} \exp(-1/2). \end{aligned} \quad (D10)$$

Meanwhile, for  $v_z \sim V$ , the first term is  $-2V/v_{th,e}^3 \exp[-(V/v_{th,e})^2]$ . Taking the ratio of these terms, we find

$$\frac{d\bar{F}_b/dv_z}{|d\bar{F}_e/dv_z|} \sim \frac{n_b}{n_e} \frac{m_e}{m_p} \frac{v_{th,e}^2}{\delta V^2} \exp[(V/v_{th,e})^2], \quad (D11)$$

where we emphasize that  $d\bar{F}_e/dv_z < 0$ . Substituting numerical values  $V = V_{DD} \approx 2.5 \times 10^9$  cm s $^{-1}$ ,  $\delta V \approx 6.3 \times 10^7$  cm s $^{-1}$ ,  $v_{th,e} \approx 8.4 \times 10^8$  cm s $^{-1}$ ,  $n_b \approx 10^8$  cm $^{-3}$ , and  $n_e \approx 10^{20}$  cm $^{-3}$ , it is clear that

$$\frac{d\bar{F}_b/dv_z}{|d\bar{F}_e/dv_z|} \ll 1, \quad (D12)$$

a consequence of the 12 orders of magnitude difference between the plasma and beam densities. Clearly, the slope of the proton-beam distribution function is too small to overcome the negative slope of the electron distribution function. Thus,

we conclude that the beam density is too low to drive electrostatic kinetic beam instabilities.

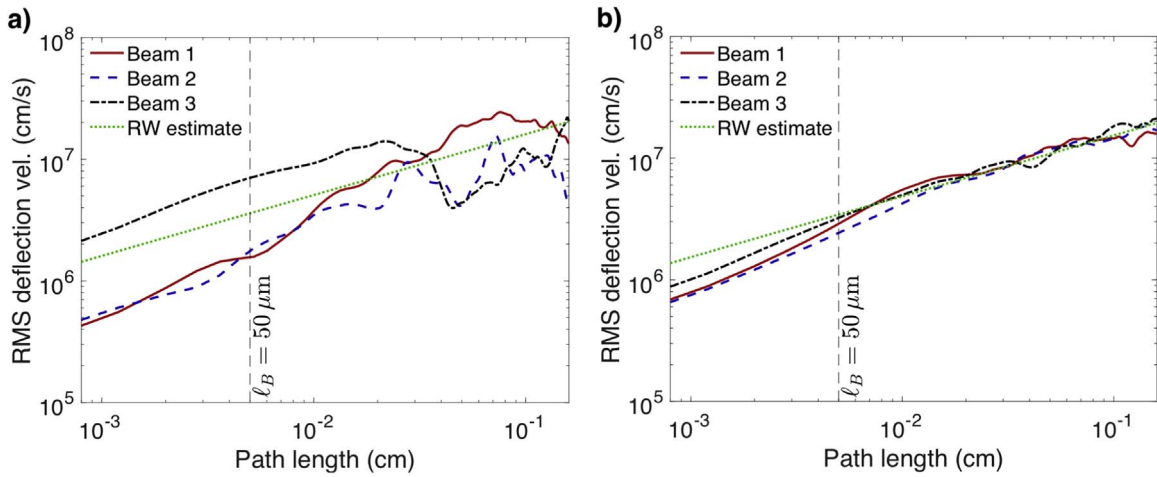
In conjunction with this conclusion, we do not expect scattering by unmagnetized plasma waves (i.e., via Landau damping) or any other electrostatic effects to be responsible for the rise in diffusion with time. First, the density and/or temperature variation from 32 to 42 ns is quite small (i.e., the Debye length is roughly constant), whereas such an effect would require a significant increase in temperature and decrease in density to promote additional scattering. Second, the electric fields to support such a wave amplification would be a significant fraction of the total energy of the system, implying an electrostatic instability at play, which, in fact, Landau damping would suppress in a plasma with near Maxwellian particle species distributions such as those relevant for our experiment.

## Appendix E

### Determination of the Range of Physical Conditions for which Diffusion is not Affected by Spatial Intermittency

As stated in Section 4, the FLASH simulations and our experimental results demonstrate that the diffusion of charged particles in the large- $r_g/\ell_B$  regime is not affected by the spatial intermittency of the stochastic magnetic fields for the physical conditions we studied. Here we describe numerical simulations we conducted to determine the range of physical conditions for which diffusion is not affected by spatial intermittency. In the simulations, test particle beams propagate through a given stochastic magnetic field, measuring the rms deflection velocity as a function of path length.

For simplicity, we do not use the FLASH-simulated fields, but instead consider homogeneous single-scale stochastic fields with a prescribed magnetic energy spectrum, chosen so that the correlation length  $\ell_B$  and rms magnetic field strength  $B_{rms}$  are the same as those found in our experiment. The results show that for a beam with a cross-sectional area  $A$  much smaller than the characteristic cross-sectional areas of magnetic structures (i.e.,  $A/4\ell_B^2 \ll 1$ )—which is effectively analogous in our parameter regime to that of a single particle—the three beams roughly obey Equation (8) for path lengths  $\ell_i$  greater than the magnetic field coherence length  $\ell_B$ , but there are significant deviations from the exact result as a function of path length, and beams with distinct initial locations have different behaviors. See, e.g., Figure E1(a). In contrast, beams with cross-sectional areas satisfying  $A/4\ell_B^2 > 1$  converge to Equation (8) for  $\ell_i > \ell_B$ , as expected. See, e.g., Figure E1(b). These results show that the use of standard diffusion theory to model the transport of charged particles applies when the magnetic field is weak, turbulent, and spatially intermittent;  $r_g \gg \ell_B$ ; and the cross-sectional area of the beam  $A$  and the path length  $\ell_i$  are greater than the magnetic field coherence length  $\ell_B$ . The transport of ultrahigh-energy cosmic rays in the intergalactic medium satisfies these conditions, so our experiment validates the use of standard diffusion theory to model it, e.g., Kotera & Lemoine (2008), Globus et al. (2008, 2019), and Globus & Piran (2017).



**Figure E1.** Dependence of the rms deflection velocity on path length and beam cross-sectional area. (a) rms deflection velocity  $\Delta v_{\perp}$  acquired by beams of 15.0 MeV protons (total number  $5 \times 10^5$ ) with the cross-sectional area  $A = 1 \times 10^{-5} \text{ cm}^2$  passing through a prescribed stochastic magnetic field as a function of path length. The full stochastic magnetic field—which has non-Gaussian exponential statistics, as well as  $B_{\text{rms}} = 80 \text{ kG}$ , and  $\ell_B = 50 \mu\text{m}$ —is generated using the approach described in Bott et al. (2017) on a  $201^3$  three-dimensional array, with (physical) side length  $\ell_i = 0.16 \text{ cm}$ . To calculate  $\Delta v_{\perp}$  at a given path length, the proton beam is propagated using the small-deflection approximation from one side through the whole array, with the rms deflection velocity recorded as the particle propagates. The initial velocity is set to be exactly perpendicular to the side on which the beam is incident. Three different beams are employed (beam 1, 2, and 3), with randomly chosen centroids. The random-walk (RW) estimate is calculated using the known properties of the field and Equation (8). (b) Same as (a), except that the beams are given cross-sectional areas  $A = 9 \times 10^{-4} \text{ cm}^2$ . For  $\ell_i > \ell_B$ , we see the expected convergence to Equation (8).

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